

# The AGN luminosity fraction in merging galaxies

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## ABSTRACT

Galaxy mergers are key events in galaxy evolution, often causing massive starbursts and fueling active galactic nuclei (AGNs). In these highly dynamic systems, it is not yet precisely known how much starbursts and AGNs, respectively, contribute to the total luminosity, at what interaction stages they occur, and how long they persist. Here we estimate the fraction of the bolometric infrared (IR) luminosity that can be attributed to AGNs by measuring and modeling the full ultraviolet to far-IR spectral energy distributions (SEDs) in up to 33 broad bands for 24 merging galaxies with the Code for Investigating Galaxy Emission. In addition to a sample of 12 confirmed AGNs in late-stage mergers, found in the *Infrared Astronomical Satellite* Revised Faint Source Catalog, our sample includes a comparison sample of 12 galaxy mergers from the *Spitzer* Interacting Galaxies Survey, mostly early stage. We also perform SED modeling of merger simulations to validate our methods, and we supplement the SEDs with mid-IR spectra of diagnostic lines obtained with *Spitzer*'s InfraRed Spectrograph. The estimated AGN contributions to the IR luminosities vary from system to system from 0 per cent up to  $\sim 91$  per cent but are significantly greater in the later-stage, more luminous mergers, consistent with what is known about galaxy evolution and AGN triggering.

**Key words:** galaxies: interactions – galaxies: photometry – galaxies: star formation – infrared: galaxies – galaxies: active – galaxies: nuclei.

## 1 INTRODUCTION

Galaxy interactions have long been known to influence galaxy evolution, and a large majority of galaxies in the universe show signs of previous interactions (e.g. Struck 1999). Mergers trigger enhanced star formation (SF) and generate or fuel active galactic nuclei (AGNs; Keel et al. 1985; Sanders et al. 1988; Hong et al. 2015; Brasington et al. 2015, and references therein). In addition, mergers produce disturbed morphological features (Toomre & Toomre 1972; Lanz et al. 2013 [hereafter L13] and references therein). Tidal

tails and nuclear disruptions are the most obvious indications that two or more galaxies are interacting or merging.

The enhanced SF seen in galaxy mergers is, in most cases, the main power source for (ultra) luminous infrared galaxies ([U]LIRGs) in the local universe (Sanders & Mirabel 1996). Veilleux, Kim & Sanders (2002) have shown that many of these galaxies contain clear morphological indicators of past galaxy interactions. However, not all galaxy mergers have enhanced IR emission. The star formation rate (SFR) depends on the merger stage, as SF and AGN activity are enhanced in the later stages of mergers (Lackner et al. 2014).

In efforts to address when and how SF and AGN activity proceed, (U)LIRGs and other luminous late-stage mergers have become prime targets for space-based telescopes such as the *Spitzer Space Telescope* (Werner et al. 2004), the *Herschel Space*

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Observatory (Pilbratt et al. 2010), the *Galaxy Evolution Explorer* (GALEX) (Martin et al. 2005), and *Swift* (Gehrels et al. 2004) among others. This suite of instruments provides highly reliable photometry by virtue of their privileged vantage point above the Earth’s atmosphere, and they are sensitive to the multiple processes contributing to galaxies’ spectral energy distributions (SEDs): *Spitzer*’s infrared (IR) cameras measure the enhanced mid-infrared (MIR) emission from both AGN and SF, *Herschel* views emission principally from the diffuse galactic dust, and the UV satellites are sensitive to the emission arising from young, hot stars.

SEDs that span the range from UV to FIR wavelengths reflect all significant energetic processes occurring in galaxies (Groves et al. 2008). They are therefore indispensable for inferring galaxies’ underlying physical properties, including but by no means limited to SFRs, masses, and luminosities of the galactic dust and the effects of AGNs (Hayward & Smith 2015). For galaxy mergers, the MIR emission arises from dust heated by both SF and AGN (Gruppioni et al. 2008). However, the relative proportions each process contributes are not well understood and vary enormously over time and from one system to another. Moreover, the ability to detect high-redshift galaxies is increasing, but the SEDs of these distant sources are much easier to obtain than spatial morphology and are therefore our best keys to understand the physical processes underway. A clear understanding of SEDs in the local universe is a prerequisite for drawing reliable conclusions about distant sources.

Many studies are being performed to calculate the fraction of luminosity that comes from the AGN in merging galaxies. Studies utilize wide wavelength ranges (from total IR to UV/X-ray) and span redshifts from the local universe ( $z < 0.06$ ) out to the distant universe ( $1 < z < 4$ ) (e.g. Ciesla et al. 2015; Drouart et al. 2016; Fernandez-Ontiveros et al. 2016; Vaddi et al. 2016; Villforth et al. 2017). Other recent studies have also characterized the SEDs out to 24  $\mu\text{m}$  of radio-loud AGNs (Williams et al. 2017) or specifically focused on the far-IR SED shape, where dust and AGN emission dominate (Safarzadeh et al. 2016; Cowley et al. 2017). Accurate constraints on the AGN contribution to the total luminosity are necessary for precise estimates of other energetic processes such as SFR. In galaxies at cosmological distances for which primarily photometric data are available, we can ascertain which physical processes are providing the luminosity for the galaxy.

This work presents an analysis of 24 merging galaxies organized into two samples. First we re-measure in a uniform and self-consistent manner and then analyze the full SEDs of 12 late-stage merging (U)LIRGs and estimate their fractional AGN contributions across the entire IR range from 1 to 1000  $\mu\text{m}$  (hereafter referred to as ‘total IR’). Our SEDs cover UV to far-IR/submillimeter wavelengths, providing a comprehensive view into the processes at work in merging galaxies. We compare the results for late-stage mergers to those for 12 previously studied reference mergers, primarily in earlier stages.

This paper is organized as follows. Section 2 presents the galaxy samples, the observations, and the data reduction. Section 3 describes the SED analysis. Section 4 describes the same type of SED analysis of merger simulations. Section 5 discusses results, and Section 6 summarizes conclusions.

## 2 OBSERVATIONS

### 2.1 Galaxy samples

We chose the late-stage merger sample to represent strongly interacting, IR-luminous systems. Candidate systems were identified by

the *Infrared Astronomical Satellite* (IRAS; Neugebauer et al. 1984) Revised Faint Source Catalogue (FSC; Wang et al. 2014) and the IRAS Revised Bright Galaxy Sample (Sanders et al. 2003). We selected interacting systems by cross-referencing IRAS sources with Version 2 of the Galaxy Zoo public galaxy classification program (Willett et al. 2013) to yield 453 systems. Of these, 85 have far-IR luminosities at least in the ‘luminous infrared galaxy’ (LIRG:  $L_{\text{IR}} > 10^{11} L_{\odot}$ ) range, and 7 are in the ULIRG ( $L_{\text{IR}} > 10^{12} L_{\odot}$ ) range. We classified all these systems by merger stage (Weiner et al. in preparation) finding 38 with morphological evidence for strong interaction such as long tidal tails or heavily disturbed morphology. These systems are designated as ‘Stage 4’ or later by Weiner et al. (in preparation). Of these 38 galaxies, only 12 had available photometry from all of *Spitzer*/IRAC, *Spitzer*/MIPS 24  $\mu\text{m}$ , and *Herschel*/SPIRE at 250, 350, and 500  $\mu\text{m}$ , respectively. These 12 constitute the late-stage merger sample listed in Table 1. By construction the sample is pure for strong interactions and high luminosities – indeed 11 of the 12 are in the top 20 luminosities of all 453 in the Galaxy Zoo sample. However, the sample is far from complete because of our requirement to have full data sets, especially *Herschel* data. This may have introduced a bias toward ‘interesting’ systems and therefore ones with extreme properties, but any bias is probably related to obvious properties such as morphology and luminosity rather than parameters that can be revealed only by detailed analysis.

For a control set to compare with the late-stage mergers, a ‘Reference Sample’ was drawn from the SIGS galaxy sample (Brassington et al. 2015). SIGS consists of 103 galaxies in 48 systems selected by a combination of galaxy proximity on the sky and morphological disturbance. SIGS therefore includes all merger stages from non-interacting systems to early approach to late stages (Brassington et al. 2015, L13). From the SIGS sample, we selected 12 galaxies with UV–submm photometry comparable to what was available for the Late-Stage Merger Sample. We adopted the merger stage classifications from L13 for these objects. Nearly all of them are Stages 2–3 implying at least mild but at most moderate distortions and galaxies still separated from each other (Weiner et al. in preparation). The Reference Sample members are listed in Table 1. The sample is deliberately heterogeneous but contains a range of systems that are merging but have not yet reached the final merger stage. The requirement for many-band photometry introduces a bias toward well-studied systems, which are likely if anything to be those with especially strong merger signatures, i.e. the Reference Sample probably resembles the Late-Stage Merger Sample more closely than the full SIGS sample would.

One difference between the samples is their redshift distributions. The Reference Sample galaxies all lie within  $z \leq 0.012$ , but the Late-Stage Merger Sample galaxies are in the range  $0.01 \leq z \leq 0.06$ . This reflects the fact that late-stage mergers are less common than early-stage ones – the early stages last longer than late stages – and it’s necessary to search larger distances to find equal numbers of late-stage systems. None of the galaxies requires a significant  $k$ -correction, and all inhabit the present-day universe (Hopkins, Richards, and Hernquist 2007).

### 2.2 Photometry

For most galaxies in the Reference Sample, we assembled photometry for single galaxies rather than the entire merging systems to ensure the most reliable SED fits. This included even the merging systems M51A/B, M101/NGC 5474, NGC 3031/3077, and NGC 5394/5395. The advanced merger NGC 3690/IC 694 system is an

**Table 1.** Basic data for the two galaxy samples, with redshifts from the SIMBAD database (Wenger et al. 2002).

Galaxy name	Redshift ( $z$ )	$D_L$ (Mpc) <sup>a</sup>	$\log L_{IR}$ ( $M_\odot$ ) <sup>a</sup>	Stage <sup>b</sup>	AGN <sup>c</sup>	Ref <sup>d</sup>
Late-Stage Merger Sample						
IRAS 08572+3915	0.05835	265	12.08	4	LINER	(1)
IRAS 15250+3609	0.05516	249	12.04	5	LINER	(1)
Mrk 231	0.04217	188	12.51	6	Seyfert 1	(4)
Mrk 273	0.03778	168	12.05	5	Seyfert 2	(2)
Mrk 463	0.05035	227	11.73	4.5	Seyfert 2	(3)
NGC 2623	0.01851	81	11.33	5	LINER	(5)
NGC 3758	0.02985	132	11.00	4.5	Seyfert 1	(2)
NGC 6090	0.02930	129	11.49	4.5	Seyfert	(2)
UGC 4881	0.03930	175	11.60	4	LINER Composite	(1,2)
UGC 5101	0.03937	175	12.03	5	Seyfert 1	(3)
VV 283	0.03748	167	11.46	5	Seyfert 2	(2)
VV 705	0.04019	179	11.82	4.5	Composite	(2)
Reference Sample						
M51A	0.00155	8.58 <sup>e</sup>	10.51	3	Seyfert 2	(6)
M51B	0.00191	8.58 <sup>e</sup>	9.61	3	LINER	(6)
NGC 2976	0.00004	3.55 <sup>f</sup>	8.91	2	None	(8)
NGC 3031	−0.00014	3.5 <sup>f</sup>	9.62	2	Seyfert 2	(3)
NGC 3077	0.00004	3.83 <sup>f</sup>	8.85	2	None	(6)
NGC 3190	0.00432	19.3 <sup>f</sup>	9.82	3	LINER	(8)
NGC 3690	0.0100	43	11.76	4 <sup>g</sup>	LINER Composite	(1)
NGC 4625	0.00212	11.75 <sup>e</sup>	8.95	3	Seyfert	(9)
NGC 5394	0.01153	50	10.69	3.5	Composite	(2)
NGC 5395	0.01158	50	10.76	3.5	Seyfert 2	(9)
M101	0.00081	6.7 <sup>f</sup>	10.30	3	None	(7)
NGC 5474	0.00098	6.8 <sup>f</sup>	8.60	3	None	(7)

Notes. <sup>a</sup>Luminosity distance  $D_L$  for the galaxies at  $z > 0.01$  were calculated using the Hubble law with  $H_0 = 70$  and scaling by  $(1 + z)$ . Total IR luminosity is calculated from 5–1000  $\mu\text{m}$ , following Fritz et al. (2006).

<sup>b</sup>Weiner et al. (in preparation).

<sup>c</sup>‘Composite’ indicates strong AGN and SF activity coexist.

<sup>d</sup>References for AGN classification: (1) Vardoulaki et al. (2015); (2) Toba et al. (2013); (3) Baumgartner et al. (2013); (4) Ivanov et al. (2000); (5) García-Lorenzo et al. (2015); (6) Hernández-García et al. (2016); (7) Brasington et al. (2015); (8) González-Martín et al. (2015); and (9) Véron-Cetty & Véron (2010).

<sup>e</sup>McQuinn et al. (2017).

<sup>f</sup>Dale et al. (2017).

<sup>g</sup>NGC 3690abc consists of two blended objects (a and b), with a nearby but separate third component c (IC 694). The tabulated photometry for NGC 3690 comes from the ab components only.

exception; it had to be observed as a single blended entity. Galaxies in the Late-Stage Merger Sample could only be imaged as a single source. Table 1 lists the physical parameters and previously known AGN status for each galaxy. Details about the SED fit for individual objects can be found in Appendix A.

Photometry for all galaxies in both samples used matched apertures on archival broadband images of up to 33 different bands (following L13). We started with the near- and far-UV bands from GALEX (Martin et al. 2005). At visible wavelengths, we used *ugriz* imaging from the Sloan Digital Sky Survey (SDSS) Data Release 12 (Alam et al. 2015). For the near-IR bands, we used *JHK<sub>s</sub>* imaging from the Two Micron All-Sky Survey (2MASS; Skrutskie et al. 2006). Mid-IR comprised *Spitzer* Infrared Array Camera (IRAC) 3.6, 4.5, 5.8, and 8.0  $\mu\text{m}$  imaging and Multiband Imaging Photometer for *Spitzer* (MIPS) 24, 70, and 160  $\mu\text{m}$  imaging. We also used the *IRAS* (Neugebauer et al. 1984) bands at 12, 25, 60, and 100  $\mu\text{m}$ . For the late-stage mergers we also incorporated near- and mid-IR imaging from the *Wide-Field Infrared Survey Explorer* (WISE; Wright et al. 2010) at 3.4, 4.6, 12, and 22  $\mu\text{m}$ . Most of the WISE imaging was taken before the W4 filter was recalibrated (Brown, Jarrett & Cluver 2014), but because the WISE 22  $\mu\text{m}$  data are outweighed in the fitting by the MIPS

24  $\mu\text{m}$  and the IRAS 25  $\mu\text{m}$  data, the difference in the SED fit between the previous and new calibrations of W4 is negligible. Finally, for the far-IR bands we used archival imaging from the *Herschel* Photoconductor Array Camera and Spectrometer (PACS) 60–90  $\mu\text{m}$ , 90–130  $\mu\text{m}$ , and 130–210  $\mu\text{m}$  bands as well as the *Herschel* Spectral and Photometric Imaging Receiver (SPIRE) bands at 250, 350, and 500  $\mu\text{m}$ . For the *Herschel* imaging we used the *Herschel* Interactive Processing Environment (HIPE), version 14.1 (Ott 2010).

Some data sets required special handling. The publicly available archival IRAC mosaics for IRAS 08572+3915 and Mrk 231 were not suitable for photometry. The Mrk 231 mosaics (specifically, the post-basic calibrated or PBCD mosaics) were saturated in all four IRAC bands. The 5.8 and 8.0  $\mu\text{m}$  PBCD mosaics for IRAS 08572+3915 also show conspicuous saturation. We therefore constructed our own mosaics for these two objects using only the short exposures (0.6 s) from archived IRAC high-dynamic range observations. After first verifying that the resulting short-exposure mosaics showed no saturation, we used them for our photometry in place of the publicly available IRAC mosaics. In addition, we adopted the global *IRAS* photometry from the *IRAS* FSC (Moshir, Kopman, and Conrow 1992).

We assembled the non-*IRAS* photometry for all 24 galaxies following the procedure described by Weiner et al. (in preparation). We started by subtracting the sky background using the `PYTHON` package *photutils*<sup>1</sup> and used an elliptical aperture to capture all the flux. We took care to use the same aperture area to enclose the full galaxy emission regions in all bands and to correct for any background emission. Our photometric values are consistent with but more precise than results in the open literature typically obtained in pipeline processing of larger data sets. The photometry can be found in Appendix B.

For the early-stage mergers from L13, we used the photometry stated in the paper but discarded the *UBV* photometry from the updated Third Reference Catalog (RC3; de Vaucouleurs et al. 1991, Corwin, Buta, and de Vaucouleurs 1994) and added the more recent SDSS *ugriz* photometry, which was processed the same as by Weiner et al. (in preparation). Photometric uncertainties were calculated using the sum in quadrature of Poisson photon noise determined by the photometry and calibration uncertainties, and we adopted the same calibration uncertainties as L13, following the references therein.

### 2.2.1 PACS spectrophotometry

We supplemented the SEDs for the late-stage merger sample with a previously underutilized resource: spectrophotometric continuum measurements taken from PACS spectral scans. The PACS spectrophotometric data provide excellent coverage of the peak of the FIR continua. Multiple observers obtained PACS range or spectral scans of lines of these sources, including some galaxies lacking standard PACS photometry. Some galaxies were observed many times, and for them the PACS archive has an abundance of spectrophotometry, while others were observed in only a few lines.

The PACS observers used a variety of observing configurations (e.g. chopping throw, integration times, scan lengths, the number of repeats), so all the PACS spectrophotometry had to be reduced individually. We used HIPE 15 and pipeline processing 14.2, which were the most recent versions of each separate program at data collection time. The sources here are adequately contained within the central  $3 \times 3$  spaxels of the IFU (a practical limit being a diameter of about 15 arcsec), and we used flux density values obtained with the C129 calibration, taking the sum of the central  $3 \times 3$  spaxels. The task used is `extractCentralSpectrum` for the `chopNod` Astronomer's Observing Templates (AOTs). We obtained the continuum level as the median of flux density values away from the spectral feature. Because each scan typically has many bad values at the start and end, those were also excluded. As a consistency check, we performed both automatic and manual flux density measurements, and they were in excellent agreement.

The PACS spectrophotometry required creating custom single-pass filters. These were  $0.5 \mu\text{m}$  wide, which corresponds to the width of the bins generated when doing off-line spectrophotometry. These filters consisted of a delta function throughput at the measured continuum wavelength. In general, we took the continuum data closest to 60, 90, and  $150 \mu\text{m}$  containing the most individual observations, such as the continuum near 63 microns or 88 microns. This allowed for the most consistent derived photometric values while also remaining near but not on top of the PACS photometry wavelengths.

There are usually hundreds of datapoints in a typical PACS spectral scan, but they are only quasi-independent. Although the formal flux density uncertainties are small because of the large number of points, we adopted a value of  $\pm 10$  per cent as more fairly accounting for the systematic uncertainties, similar to the PACS photometric uncertainties (Paladini et al. 2012). See Appendix B for the table of PACS spectrophotometry for the 10 galaxies in the late-stage merger sample containing PACS spectra.

We did not obtain PACS spectrophotometric data for the early-stage merging galaxies from the L13 sample because of their low modeled  $f_{\text{AGN}}$  from CIGALE. The galaxies in the late-stage sample all had large  $f_{\text{AGN}}$ , and the AGN emission models from Fritz, Franceschini & Hatziminaoglou (2006) peak at the PACS wavelength range, so we wanted to have the best characterization possible for the emission in the PACS range for the late-stage merging galaxies with high AGN fractions. However, because this adds multiple new data points in a small wavelength range, the risk of overfitting the SEDs increases. For the galaxies with large AGN fractions and a high sensitivity to changes in the PACS bands, only three spectrophotometric values were used with similar uncertainties to the PACS photometric data to help characterize the peak without overfitting. To avoid the issue for the galaxies with low AGN fractions in the early-stage sample, where the emission peak is already well characterized, we omitted PACS spectrophotometric data.

### 2.3 Spectroscopy

Another way to estimate the AGN luminosity fraction is by using spectral lines that separately trace AGN and SF activity. Specifically, [O IV] at  $25.89 \mu\text{m}$  and [Ne V] at  $14.32 \mu\text{m}$  are strong lines that signify the presence of an AGN. In contrast, the [Ne II] line at  $12.81 \mu\text{m}$  is diagnostic of SF activity (Dale et al. 2009; LaMassa et al. 2012). We used spectra taken by *Spitzer*'s InfraRed Spectrograph (IRS; 2004 Houck et al.), which provided spectroscopic coverage from  $5\text{--}36 \mu\text{m}$ . For each galaxy, we used the Short-High (SH) and Long-High (LH) modes, which have resolving power  $\lambda/\Delta\lambda = 600$ . The SH mode uses a slit of  $4.7 \times 11.3$  arcsec, while the LH mode uses a slit of  $11.1 \times 22.3$  arcsec.

We reduced IRS spectra using the Spectroscopic Modeling Analysis and Reduction Tool (SMART; Higdon et al. 2004), version 8.2.9. SMART was specifically designed for the IRS and provides an easy-to-use interface to reduce, analyze, and view the spectra. To validate our results from SMART, we compared our final IRS spectra for a few galaxies with those reduced using an earlier version of SMART and placed on the Cornell Atlas of *Spitzer*/IRS Sources (CASSIS; Lebouteiller et al. 2011). Differences in the two versions of the final spectra were negligible. We used the data stored on CASSIS for spectral line analysis for 14 galaxies in our sample because the spectra reduction was already completed and reliable. For the remaining 10 galaxies, we used the results from SMART.

For these 10 galaxies, we retrieved the IRS spectra from the *Spitzer* Heritage Archive. The observations were taken in both Stare and Map modes. We extracted Stare observations flux densities and spectra without further corrections, but for the Map images we used the Cube Builder for IRS Spectral Mapping (CUBISM; Smith et al. 2007) to build 3D spectral cubes with 2 spatial and 1 spectral dimensions. The 2 spatial dimensions had pointings in a  $3 \times 3$  grid with the centre pointing aimed at the nucleus. We confirmed the central spectra contained the galaxy nuclei before extracting the results to files for importing into SMART. Then, we used SMART to fit and extract the three target spectral line features.

<sup>1</sup><https://github.com/astropy/photutils>



To start, we calculated a linear baseline around each spectral line and subtracted it before fitting a Gaussian to the line profile. For the cases where the signal was not strong or significant contamination caused irreversible blending, SMART would provide a Gaussian fit with limits, so these translated into  $3\sigma$  upper limits on the integrated line flux and the line width.

The [Ne II] line is relatively isolated with no other nearby, potentially contaminating spectral lines. The [Cl II] line at  $14.37 \mu\text{m}$  caused no noticeable contamination to the [Ne V] line. However, the [O IV]  $25.89 \mu\text{m}$  line partially overlaps with the [Fe II] line at  $25.99 \mu\text{m}$ , producing a slightly blended line profile. For lines with significant potential contamination, we used a double Gaussian profile to fit the composite (double-line) profiles. The [Fe II] contamination did end up forcing the use of an upper limit for some detections of the [O IV] line as the blending caused both single and double Gaussian fits to fail.

We calculated integrated line fluxes and widths from the Gaussian fits. We compared the results for [Ne V]/[Ne II] and [O IV]/[Ne II], similar to the analysis of Genzel et al. (1998). Because [Ne II] is a strong tracer of starburst activity, whereas [Ne V] and [O IV] are strong indicators of AGN activity, the AGN-to-starburst tracer ratios help determine the dominant source of luminosity for these galaxies (Armus et al. 2007; Satyapal et al. 2009 and references therein). Higher ratios should indicate larger AGN activity compared to starbursts. Ramos Padilla et al. (in preparation) provide a detailed analysis of more spectral line ratios and their correlations with IR colours that indicate the presence of AGNs.

### 3 ANALYSIS

#### 3.1 SED fitting

For the SED fitting, we used the Code for Investigating Galaxy Emission (CIGALE; Burgarella, Buat & Iglesias-Paramo 2005). Specifically, we used ‘PCIGALE’ version 0.9.0 in PYTHON, which was released in 2016 April. In brief, CIGALE operates by constructing a multidimensional grid of model SEDs and identifying the SED model that best fits the data with  $\chi^2$  minimization. The grid dimension is set by the number of user-defined parameters used to define the different galaxy components, e.g. intrinsic AGN and stellar emission spectra, star formation history, dust attenuation, and nebular emission. After it has tested all user-specified models in its grid, CIGALE then outputs what it identifies as the best-fitting model spectrum and the parameter set that best matches the galaxy data. CIGALE also outputs parameter uncertainties based on the range of models that are consistent within each galaxy’s flux density uncertainties.

In this work, for simplicity we used a ‘delayed’ star formation history model (delayed with respect to the SF timescale), assuming a single starburst with an exponential decay, following

$$\text{SFR}(t) \propto \frac{te^{-t/\tau}}{\tau^2}, \quad (1)$$

where  $\tau$  is the e-folding time of the main stellar population, which dominates the stellar emission (Lee et al. 2010). The SF starts at a time ‘age’ before the present day, where ‘age’ is a CIGALE parameter given in the model and defined in Table 2 (Ciesla et al. 2016). We also set the separation between the young and old stellar populations (the stellar separation age) to 10 Myr. This means that at the time that the galaxy is modeled in CIGALE, every star older than 10 Myr is considered ‘old’ while the rest are considered ‘young’. The combination between  $\tau$  (defined in CIGALE as  $\tau_{\text{main}}$ ), ‘age’,

and stellar separation is used as a proxy to model the recent star formation in the period of time defined by ‘stellar separation’. This parametric SFH model allows for CIGALE to be tuned to the recent SFR and can help determine the stage in some complicated cases. Tests running other SF history options (e.g., Leitherer et al. 1999) did not significantly alter our conclusions about the AGN fraction.

For the dust attenuation, we used models jointly described by Calzetti et al. (2000) and Leitherer, Calzetti & Martins (2002) along with the Dale et al. (2014) models for the dust emission in the far-IR. The Calzetti law for dust extinction and attenuation is described by the following set of local piece-wise power-laws,

$$\kappa(\lambda) = \frac{A(\lambda)}{E(B-V)_*} = a + \frac{b}{\lambda} + \frac{c}{\lambda^2} + \frac{d}{\lambda^3}, \quad (2)$$

where  $a$ ,  $b$ ,  $c$ ,  $d$  are constants dependent on the wavelength range. The dust emission from Dale et al. (2014) follows a modified black-body SED with a power-law distribution of dust mass at each temperature,

$$dM_d \propto U^{-\alpha} dU, \quad (3)$$

where  $M_d$  is the dust mass heated by a radiation field at intensity  $U$ . The power-law index  $\alpha$  was allowed to vary from 1 to 3. We used the stellar emission models from Bruzual & Charlot (2003) and the standard default nebular emission model included in CIGALE.

For the AGN emission we used the Fritz et al. (2006) AGN emission models, which assume isotropic emission from a central source and emission from a surrounding toroidal dust structure. The assumed central point-like luminous source was defined with a composite power-law in  $\lambda L(\lambda)$ . In particular, from  $0.001$  to  $0.03 \mu\text{m}$ ,  $\lambda L(\lambda) \propto \lambda^{1.2}$ ; from  $0.03$  to  $0.125 \mu\text{m}$ , it is independent of wavelength; and from  $0.125 \mu\text{m}$  to  $20 \mu\text{m}$ ,  $\lambda L(\lambda) \propto \lambda^{-0.5}$  (Granato & Danese 1994; Nenkova, Ivezić & Elitzur 2002). The rest of the IR emission comes from the blackbody emission due to AGN heating of the torus. The AGN emission was calculated for an intermediate-type AGN with an axis angle of  $30^\circ:1$  (where  $0^\circ$  corresponds to a Seyfert 1 galaxy viewed pole-on and  $90^\circ$  corresponds to a Seyfert 2 galaxy viewed edge-on). The default  $30^\circ:1$  axis angle was provided to CIGALE at the outset, and there was no noticeable distinction between different intermediate angles when tested against the simulated SEDs. Viewing the AGN from  $30^\circ:1$  provides enough dust extinction from the torus surrounding the accretion disk to attenuate the UV emission from the AGN, only leaving IR diagnostics (Snyder et al. 2013).

We held many CIGALE parameters constant but varied in particular the parameters defining the AGN, dust attenuation, and star formation history to model these galaxies as accurately as possible. Table 2 lists the values and/or ranges for the parameters used. With these settings, every CIGALE run calculated the reduced  $\chi^2$  for each of 6.3 million model SEDs. The model SED with the lowest  $\chi^2$  was saved along with probability density functions (PDFs) for each parameter and a text file containing the best models along with the estimates and uncertainties for each parameter. These uncertainties were derived from the  $1\sigma$  standard deviations of the PDFs created by CIGALE for each parameter. Thus the ‘best-fitting’ SED was not always from the ‘most probable’ individual parameters as found in the PDFs, but they were generally within the uncertainties, in particular for the AGN fraction.

For each best-fitting SED identified by CIGALE, we also calculated the AGN luminosity fraction (denoted ‘fracAGN’ by Fritz et al. 2006 but hereafter referred to as  $f_{\text{AGN}}$ ).  $f_{\text{AGN}}$  is defined as the AGN contribution to the total IR luminosity from  $\sim 5$ – $1000 \mu\text{m}$ . We tested values of  $f_{\text{AGN}}$  that ranged from 0 to 0.9 (90 per cent of the

**Table 2.** CIGALE parameter settings used in this work.

Parameter	Definition	Values tested (range)
Star Formation History—Delayed Module		
$\tau_{\text{main}}$	The e-folding time of the main population (Myr)	50, 500, 1000, 2500, 5000, 7500
age	The age of the oldest stars (Gyr)	0.5, 1, 2, 3, 4, 5, 6
sfr_A	Multiplicative factor controlling SFR amplitude	1.0
separation age	Separation between young and old stellar populations (Myr)	10
Bruzual & Charlot (2003) Stellar Emission Module		
imf	Initial mass function (0 for Salpeter, 1 for Chabrier)	0
metallicity	Initial metallicity for the stars	0.02
separation_age	Age of separation between ‘young’ and ‘old’ stellar populations in Myr	10
Nebular Emission Module		
logU	Ionization parameter	−2.0
f_esc	Escape fraction of Lyman continuum photons	0.0
f_dust	Absorption fraction of Lyman continuum photons	0.0
lines_width	Line width in km/s	300
emission	Whether or not to include nebular emission	True
Calzetti et al. (2000) and Leitherer et al. (2002) Dust Attenuation Module		
E_BVs_young	$E(B - V)_*$ , the colour excess of the young stellar continuum light	0.1, 0.25, 0.4, 0.55, 0.7
E_BVs_old_factor	Reduction factor for the $E(B - V)_*$ of old versus young population	0.22, 0.44, 0.66, 0.88
	Amplitude of the 220 nm bump	0.0
uv_bump_amplitude		
powerlaw_slope	Slope delta of the power law attenuation curve	0.0
filters	Filters in which attenuation will be calculated	FUV
Dale et al. (2014) Dust Module		
$\alpha$	Slope of the dust temperature distribution in 3	1, 1.5, 2, 2.5, 3
Fritz et al. (2006) AGN Module		
r_ratio	The ratio between outer and inner radius of AGN torus	10, 30, 60, 100, 150
$\tau$	The optical depth at $9.7\mu\text{m}$	0.6, 1, 6, 10
$\beta$	The density radial exponent	−1, −0.75
$\gamma$	The density exponential factor	0, 2
opening_angle	The opening angle of the torus	60, 100, 140
$\psi$	The angle between equator and line of sight (0 is Type 2 and 89.9 is Type 1)	30.1
$f_{\text{AGN}}$	The AGN fraction to the IR luminosity	0, 0.1, 0.2, 0.3, 0.4, 0.5, 0.6, 0.7, 0.8, 0.9 (0.95, 0.99)

total IR luminosity) in increments of 0.1 on all merging galaxies.  $f_{\text{AGN}} = 0$  per cent accounts for the possibility that an AGN might not contribute to the IR luminosity. Once the full grid from 0 to 0.9 was tested, CIGALE was run again on each galaxy with a finer but narrower grid for  $f_{\text{AGN}}$  centered on the best-fitting value from the previous run. Although it is theoretically possible to obtain an AGN fraction close to 1, the probability is extremely low even for the strongest AGN-dominated galaxies in our sample; empirically, we found only one case for which the AGN fraction was significantly larger than 85 per cent. For IRAS 08572+3915,  $f_{\text{AGN}}$  was allowed to exceed 90 per cent in the model runs with ranges from 0.7 to 0.95 by 0.05 along with 0.99. CIGALE found a best-fitting  $f_{\text{AGN}}$  value of 91 per cent, which was obtained through interpolation of the parameter grid points to find the best-fitting solution as measured by the reduced  $\chi^2$  value. Table 3 contains the best-fitting  $f_{\text{AGN}}$  values and reduced  $\chi^2$  for each galaxy along with the line ratios described in Section 2.3. Fig. 1 shows an example best-fitting model, Fig. 2 shows a calculated probability measure for  $f_{\text{AGN}}$ , and Fig. 3 shows  $f_{\text{AGN}}$  as a function of  $100\mu\text{m}$  luminosity for both samples. The Late-Stage Sample has larger luminosities than most of the Reference Sample by construction, but nothing selected for or against AGN fraction in either sample. If anything, AGN of a given luminosity should be easier to detect in low-luminosity galaxies, i.e. in the Reference Sample.

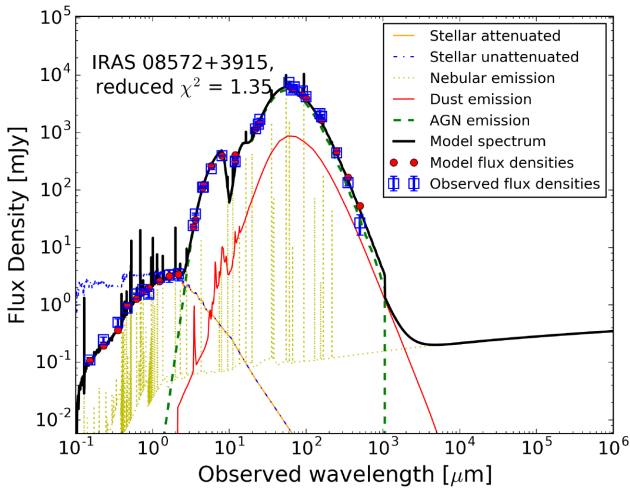
In some cases, CIGALE produced a best-fitting model with  $f_{\text{AGN}} = 0$  having no estimated uncertainty, even for galaxies known to host an AGN from previous studies. In general when  $f_{\text{AGN}} < 0.1$ , the uncertainties tend to be large fractions of the best-fitting value. As a result, the AGN model is relatively poorly constrained in low- $f_{\text{AGN}}$  cases. This is because with just photometric data, a weak AGN cannot be distinguished from a slightly increased SFR. CIGALE uses a slightly different, more luminous dust model in the IR to account for the small influx to the SED that the AGN contributes (Ciesla et al. 2015). However, in most of these cases, the CIGALE best-fitting model underestimates the far-IR *Herschel*/SPIRE bands, up to a factor of 1.5 or 2. CIGALE models including non-thermal radio emission were considered in an attempt to better fit the *Herschel*/SPIRE data points, but the radio emission was ultimately discarded as negligible because the added emission in the far-IR was 3 orders of magnitude too small to make up the difference.

For the spectral line analysis, with the results shown in Fig. 4, we ran linear regression tests on the combination of both samples for both  $[\text{Ne v}]/[\text{Ne ii}]$  versus  $f_{\text{AGN}}$  and  $[\text{O iv}]/[\text{Ne ii}]$  versus  $f_{\text{AGN}}$ , and the results were not significant at the  $3\sigma$  level. However, when running the linear regression test on the late-stage merger sample only, there was evidence for a linear trend at the  $2\sigma$  level. Further analysis of spectral line ratios, including correlations with flux density ratios, is discussed by Ramos Padilla et al. (in preparation).

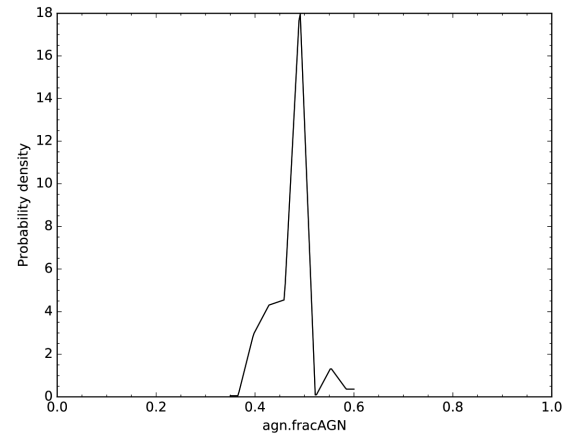
**Table 3.** Derived fractional AGN contribution to the total IR luminosity, SFR, and reduced  $\chi^2$  from the CIGALE models, and the measured [Ne v]/[Ne ii] & [O iv]/[Ne ii] ratios from *Spitzer* IRS spectra.

Galaxy name	$f_{\text{AGN}}$	SFR ( $M_{\odot} \text{ yr}^{-1}$ )	Reduced $\chi^2$	[Ne v]/[Ne ii]	[O iv]/[Ne ii]
Late-Stage Merger Sample					
IRAS 08572+3915	$0.91 \pm 0.05$	$21.4 \pm 4.9$	1.35	$<0.107$	$<0.574$
IRAS 15250+3609	$0.47 \pm 0.04$	$89.8 \pm 8.1$	2.31	$<0.101$	$<0.202$
Mrk 231	$0.17 \pm 0.02$	$444 \pm 23$	3.18	$<0.003$	$<0.362$
Mrk 273	$0.66 \pm 0.04$	$57.9 \pm 7.5$	1.66	$0.227 \pm 0.014$	$1.206 \pm 0.091$
Mrk 463	$0.68 \pm 0.03$	$30.0 \pm 2.7$	2.95	$2.038 \pm 0.136$	$6.866 \pm 0.733$
NGC 2623	$0.39 \pm 0.05$	$20.0 \pm 1.4$	3.88	$0.062 \pm 0.007$	$0.173 \pm 0.015$
NGC 3758	$0.30 \pm 0.03$	$7.02 \pm 1.12$	1.36	...	...
NGC 6090	$0.26 \pm 0.05$	$35.0 \pm 2.3$	2.00	$<0.009$	$0.045 \pm 0.021$
UGC 4881	$0.51 \pm 0.03$	$23.2 \pm 1.9$	1.39	$<0.005$	$0.037 \pm 0.008$
UGC 5101	$0.76 \pm 0.04$	$20.6 \pm 3.8$	5.56	$0.152 \pm 0.018$	$0.158 \pm 0.018$
VV 283	$0.47 \pm 0.04$	$20.7 \pm 1.9$	1.79	$0.008 \pm 0.001$	$0.015 \pm 0.004$
VV 705	$0.25 \pm 0.08$	$77.1 \pm 8.8$	0.97	$<0.006$	$<0.019$
Reference Sample					
M51A	$0.09 \pm 0.03$	$2.73 \pm 0.14$	0.96	$0.032 \pm 0.002$	$0.238 \pm 0.009$
M51B	$<0.03$	$<0.11$	0.88	$<0.057$	$<0.108$
NGC 2976	$0.28 \pm 0.06$	$0.078 \pm 0.004$	1.34	$<0.133$	$<0.041$
NGC 3031	$<0.01$	$0.382 \pm 0.019$	1.57	$<0.024$	$0.149 \pm 0.029$
NGC 3077	$0.34 \pm 0.06$	$0.069 \pm 0.014$	0.78	$<0.003$	$<0.048$
NGC 3190	$<0.01$	$0.040 \pm 0.029$	1.22	$<0.081$	$0.113 \pm 0.029$
NGC 3690	$0.30 \pm 0.05$	$59.3 \pm 3.2$	2.01	$<0.007$	$<0.027$
NGC 4625	$<0.10$	$0.105 \pm 0.005$	0.68	$<0.141$	$<0.044$
NGC 5394	$0.58 \pm 0.03$	$3.09 \pm 0.15$	2.19	$0.005 \pm 0.002$	$0.016 \pm 0.001$
NGC 5395	$<0.10$	$6.27 \pm 0.31$	0.48	$<0.021$	$<0.028$
M101	$<0.14$	$4.01 \pm 0.20$	0.68	$0.111 \pm 0.041$	...
NGC 5474	$<0.05$	$0.006 \pm 0.0003$	4.08	$<1.02$	$<0.38$

Notes. <sup>a</sup>Upper limits are defined at  $3\sigma$ . <sup>b</sup>‘...’ indicates the galaxy was not observed in the required wavelengths by IRS



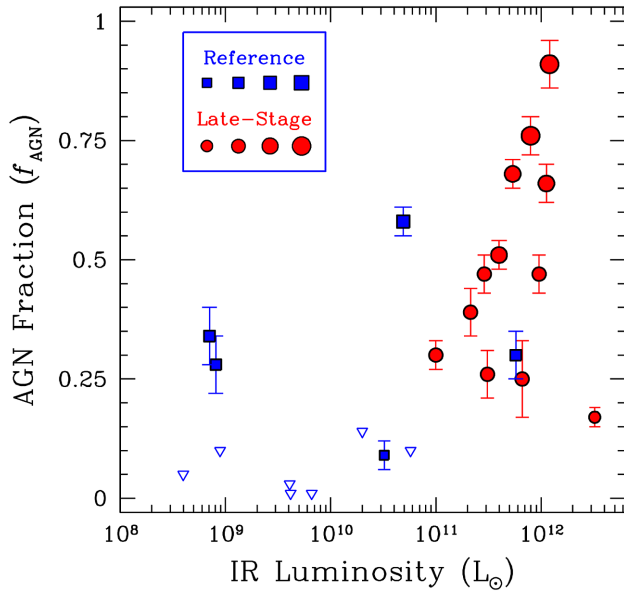
**Figure 1.** An illustration of the data quality and CIGALE SED modeling. The SED shown (blue symbols in upper panel) is for IRAS 08572+3915. The best-fitting CIGALE model is shown in black. Red dots indicate CIGALE-derived photometry in the modeled passbands. The best-fitting CIGALE model is the sum of contributions from an AGN (green dashed line), dust-attenuated stellar emission (orange; the intrinsic stellar emission is indicated in blue), nebular emission (yellow), and dust emission (red). The bottom panel shows the fractional discrepancies between the model and photometry. The best-fitting CIGALE SEDs for all 24 galaxies analyzed in this work are in Fig. A1.



**Figure 2.** A measure of the  $f_{\text{AGN}}$  probability for IRAS 15250+3609. CIGALE found 0 probability for values of  $f_{\text{AGN}}$  less than 0.3 and greater than 0.6.

### 3.2 AGN observables

Numerous studies (e.g. Stern et al. 2005; Stern et al. 2012; Donley et al. 2008; Rosario et al. 2012) have demonstrated that galaxy colours can reveal AGN. Flux ratios such as  $f_{24\mu\text{m}}/f_{100\mu\text{m}}$  and  $f_{12\mu\text{m}}/f_{24\mu\text{m}}$  as well as  $K-L$  and  $L-M$  colours have been used to help determine the presence of an AGN. We performed a linear regression test of  $f_{\text{AGN}}$  versus every flux ratio in our photometric data to determine whether any ratio showed a significant correlation. The significant results, with Pearson correlation ratios of magnitude

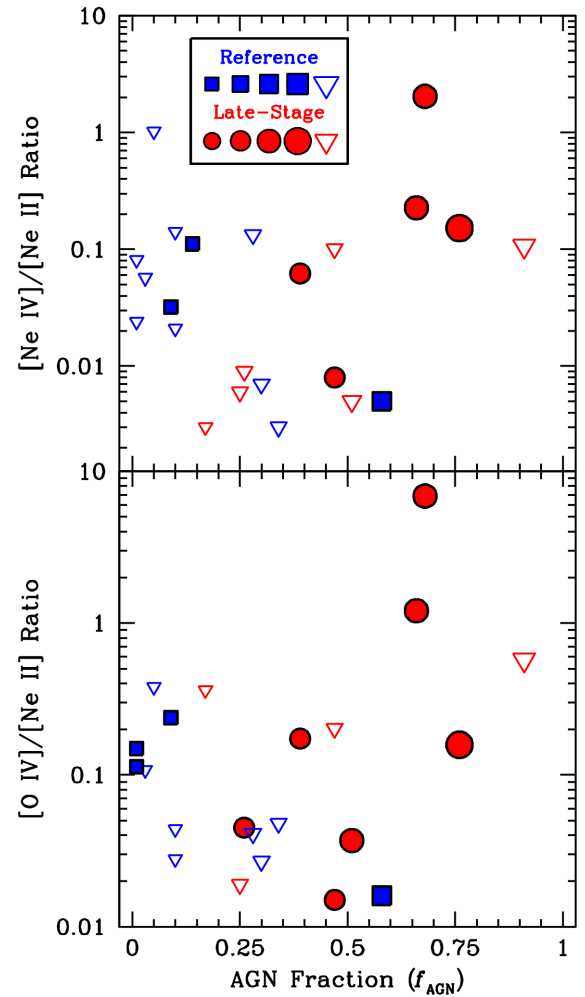


**Figure 3.** Estimated AGN luminosity fractions versus IR luminosity for both the Late-Stage Merger galaxies (red symbols) and the Reference Sample (blue symbols). The luminosity fractions were measured as a function of IR luminosity between 5 and 1000  $\mu\text{m}$ . Symbol size indicates the percentage AGN luminosity fraction: the smallest symbols are for  $f_{\text{AGN}} < 0.25$ , the next largest indicate  $0.25 \leq f_{\text{AGN}} < 0.50$ , and so on. Triangles indicate  $3\sigma$  upper limits.

greater than 0.8 and with  $p$ -values of 0.027 or less (corresponding to a significance level of  $3\sigma$ ), are summarized in Table C1.

$L - M$  and similar colours such as IRAC [3.6] – [4.5] and *WISE* W1 – W2 are the basis for the Stern et al. (2005) and Donley et al. (2008) plots showing a significant difference for the AGN-dominated systems. The [3.6] – [4.5]  $\mu\text{m}$  colour is significant here at  $\sim 6\sigma$ , as seen in Fig. 5. However, the corresponding *WISE* colours and mixing of IRAC and *WISE* colours are not significant due to the low sample size of galaxies with *WISE* photometry. Consistent with the SED analysis described in Section 3.1, the early-stage mergers consistently populate the starburst region of the Stern et al. (2005) IRAC colour–colour plot. By contrast, even though our analysis shows that all the late-stage merger galaxies contain luminous AGN and moreover that many of them are AGN-dominated, only about half of them have IRAC colours indicating that these objects host luminous AGN. The apparent discrepancy is not surprising because the SED analysis is based on much more information than the simple colour–colour plot; in particular, it attempts to estimate and account for dust obscuration. The Stern et al. (2005) plot could miss AGN when stars overwhelm the AGN at IRAC wavelengths or when the AGN is heavily obscured by dust in the near-mid IR. However, the four largest AGN fractions modeled by CIGALE correspond to the four galaxies with the reddest [3.6]–[4.5] colours. Fig. 6 contains the Stern colour–colour plot for the 24 galaxies in our sample. The galaxies from the Reference Sample nearly all fall in this region. By contrast, the late-stage merger subsample populates both the AGN wedge and the star-formation region of the plot, albeit only the extreme red end of the latter.

Sanders et al. (1988) showed that the [25] – [60] colour is an AGN tracer, but this colour is not correlated with our measurements of  $f_{\text{AGN}}$  as given by CIGALE. Two of the galaxies modeled with strong AGN are among the bluest in [25] – [60] colour, and the reddest [25] – [60] measurement belongs to a galaxy modeled with



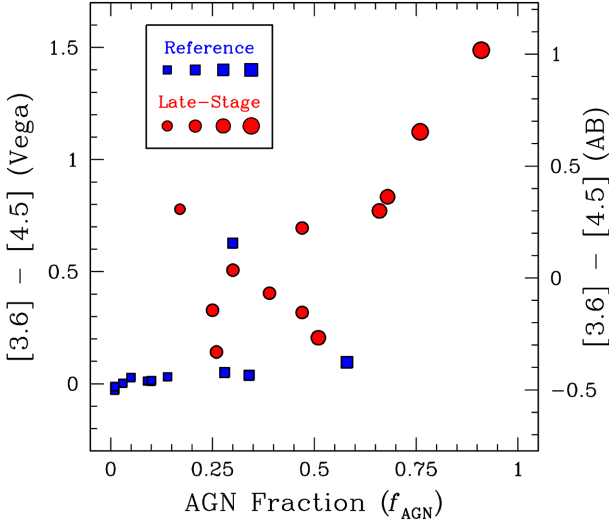
**Figure 4.** Integrated emission-line flux ratios as a function of AGN luminosity fraction. *Upper panel:* [Ne IV]/[Ne II]. Symbol size indicates the percentage AGN luminosity fraction: the smallest symbols are for  $f_{\text{AGN}} < 0.25$ , the next largest indicate  $0.25 \leq f_{\text{AGN}} < 0.50$ , and so on. Triangles indicate  $3\sigma$  upper limits. *Lower panel:* [O IV]/[Ne II].

an AGN fraction of  $\sim 5$  per cent. Ramos Padilla et al. (in preparation) analyzed the correlation this colour has with spectral line ratios in addition to AGN fraction. The UV, optical, and 2MASS near-IR connected with IRAC, IRAS, MIPS, and PACS mid and far-IR have strong correlations with  $f_{\text{AGN}}$ . In particular, the flux ratios of *GALEX* FUV and NUV, Sloan  $u$ ,  $g$ ,  $r$ ,  $i$ ,  $z$ , and 2MASS  $J$ ,  $H$ , and  $K_s$  with IRAC 4.5 and 5.8  $\mu\text{m}$  and MIPS 24  $\mu\text{m}$  and 70  $\mu\text{m}$  bands are significantly correlated with AGN fraction. The extremely red colors at the high end of the correlation show that a steep increase in the SED in the near and mid-IR is indicative of an AGN. Photometric data at similar wavelengths show similar correlations; UV–70  $\mu\text{m}$  flux ratios are significant for both MIPS and PACS 70  $\mu\text{m}$  photometry, for example. Also, negative correlations are found between the AGN fraction and the MIPS 70  $\mu\text{m}$  and PACS/SPIRE colours, showing that the presence of an AGN makes the far-IR SED steeper than the expected cool dust power-law.

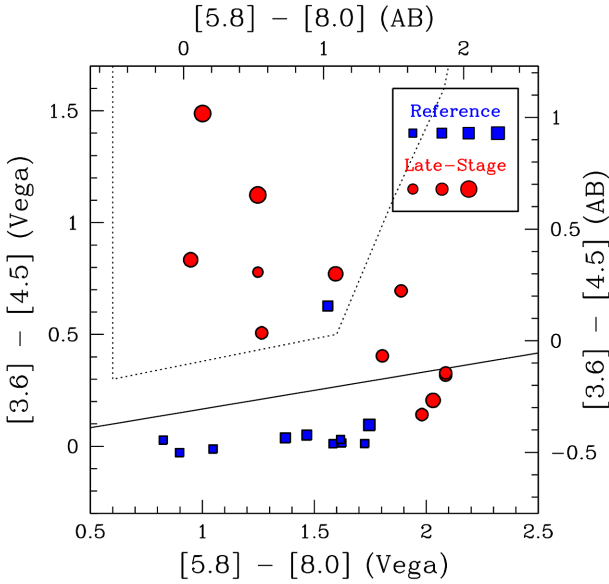
## 4 SIMULATIONS

We tested the reliability of CIGALE by analyzing the SEDs of simulated merging galaxies and comparing the CIGALE-derived





**Figure 5.** IRAC [3.6] – [4.5] color versus AGN luminosity fraction. Symbol size indicates the percentage AGN luminosity fraction: the smallest symbols are for  $f_{\text{AGN}} < 0.25$ , the next largest indicate  $0.25 \leq f_{\text{AGN}} < 0.50$ , and so on.



**Figure 6.** IRAC colour-colour plot following Stern et al. (2005) for all galaxies in both subsamples analyzed in this work. Symbol size indicates the percentage AGN luminosity fraction: the smallest symbols are for  $f_{\text{AGN}} < 0.25$ , the next largest indicate  $0.25 \leq f_{\text{AGN}} < 0.50$ , and so on. The dotted line encloses the region in which low-redshift galaxies reside when their IRAC colours are dominated by luminous AGN (the AGN wedge). The straight line is an empirical boundary below which nearby star-formation-dominated galaxies typically lie.

galaxy parameters with the known galaxy parameters from the simulations. Performing such ‘numerical experiments’ using simulations is a very useful means to validate methods of observational inference, as the ground truth is known *a priori* and various uncertainties can be controlled. For previous examples and discussions of this validation process, see Michałowski et al. (2014), Hayward & Smith (2015), Smith & Hayward (2015). As Lanz et al. (2014) have described, the simulations provide realistic SEDs (see also Weiner et al. in preparation). The aim was to determine how well

CIGALE recovers  $f_{\text{AGN}}$  (i.e. the AGN’s contribution to the bolometric luminosity) of the simulated galaxies. The simulated merger SEDs were created using a two-step process. First, mergers were simulated using a hydrodynamical code (Springel 2005; Hayward et al. 2011; Hayward et al. 2014b, Lanz et al. 2014), and then a radiative transfer code was used to generate the emergent light from the simulated mergers and simulate an observation (Jonsson 2006; Jonsson, Groves & Cox 2010). The hydrodynamic simulations and radiative transfer code used are described in detail by Lanz et al. (2014) and Weiner et al. (in preparation). Here we summarize the key aspects of our analysis.

#### 4.1 Models

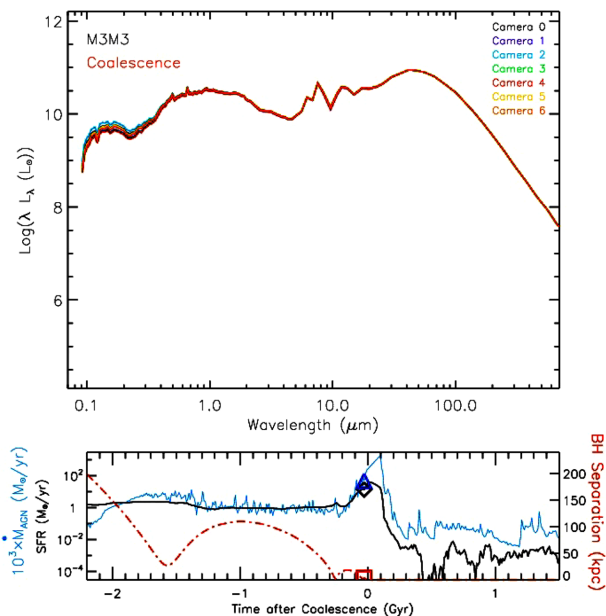
The merger simulations used the TreeSPH (Hernquist & Katz 1989, Katz, Weinberg, and Hernquist 1996) code GADGET-3 (Springel 2005), which employs a hierarchical tree N-body method to compute gravitational interactions in an N-body cosmological simulation that includes gravity, gas dynamics (Lucy 1977, via smoothed-particle hydrodynamics as in Gingold & Monaghan 1977, Springel 2010), stellar evolution, and other physical mechanisms. GADGET-3 implements the thermodynamic transport of energy through gas dynamics and radiative heating and cooling and conserves both energy and entropy. The ISM is modeled with two phases of matter in which cold, dense clouds interact with a hot, diffuse gas medium (Springel & Hernquist 2003). The hydrodynamical code models star formation according to the Kennicutt–Schmidt (‘K–S’) relation (Kennicutt 1998), an empirical relation between SFR and the volume density of gas. When the density of gas particles in the simulation surpasses a minimum threshold, gas particles are converted into star particles according to the K–S relation. GADGET-3 uses sub-resolution models to describe starforming regions because its grid is too coarse to resolve individual cold gas clouds; this limitation directly affects how radiative transfer is modeled.

The AGN contributions to the SEDs were computed from the black hole accretion rate, and the corresponding AGN feedback was included using the sub-resolution model of Springel et al. (2005). The AGNs were represented in the simulations by black hole particles which grow and radiate by accreting surrounding gas (Springel et al. 2005). Black hole particles accreted according to the Bondi-Hoyle-Lyttleton model, at the rate

$$\dot{M}_{\text{BH}} = \frac{4\pi\alpha G^2 M_{\text{BH}}^2 \rho}{(c_s^2 + v^2)^{3/2}}, \quad (4)$$

where  $\rho$  is the gas density,  $c_s$  is the speed of sound in the gas,  $v$  is the black hole speed relative to the gas, and  $\alpha$  is a system-dependent, dimensionless parameter, usually estimated as between 1 and 2 (Bondi 1952); we took  $\alpha = 1.5$ . Because the accretion occurs on spatial scales far below GADGET-3’s resolution, the code uses a sub-resolution model to interface black hole particles to surrounding gas particles. GADGET-3 models accretion of gas particles as a stochastic process. Each particle near a black hole is assigned a probability of accretion weighted by the estimated gas density near the black hole, the location of the particle relative to the BH, the Bondi accretion rate, and the time step. GADGET-3 also imposes an upper limit on  $\dot{M}_{\text{BH}}$  at the Eddington rate, at which the radiation pressure from an AGN overcomes the gravitational attraction of the gas.

As an AGN accretes gas, its accretion disk heats up and radiates energy into the host galaxy. GADGET-3 treats the thermal energy delivered to the black hole as thermal energy radiated into the



**Figure 7.** The SUNRISE SED output for the M3–M3 major merger. (Top) The full SED just prior to coalescence of the two supermassive black holes. The seven color-coded viewing angles in this case give nearly identical SEDs. (Bottom) The SFR (black), AGN accretion rate (blue) and black hole separation versus merger time in Gyr with respect to coalescence. The diamond and square markers indicate the moment corresponding to the SED in the upper panel.

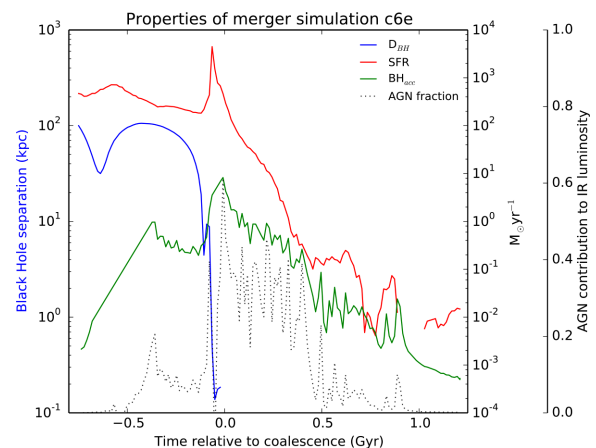
surroundings with power

$$L_r = \epsilon_r \dot{M}_{\text{BH}} c^2, \quad (5)$$

where  $\epsilon_r$  is the radiative efficiency, which is set to 10 per cent in these simulations, the consensus value for efficient black hole accretion. As can be seen in equation (5), the AGN luminosity is directly proportional to the accretion rate, so when the AGN is turned off, as described below,  $\epsilon_r = 0$ . In this way, the accreting AGN directly influence surrounding regions of star formation.

We used the 3D polychromatic Monte Carlo dust radiative transfer code SUNRISE (Jonsson 2006; Jonsson et al. 2010) to calculate spatially resolved UV–mm SEDs for the simulated galaxies. SUNRISE performs a radiative transfer calculation for the attenuation and re-emission from the dust heated by star formation and AGN activity, as well as the stellar components, to generate ‘observed’ SEDs for the merger. Merger steps for SED calculation were at regular intervals at 10 Myr near coalescence and at 100 Myr otherwise (Lanz et al. 2014). SEDs were computed for seven different viewing angles at each step to account for the impact of dust attenuation along different lines of sight.

Five galaxy models called M4, M3, M2, M1, and M0 with stellar masses respectively of 5, 4.22, 1.18, 0.38, and  $0.061 \times 10^{10} M_\odot$  were used (see Table 2 of Lanz et al. 2014; Rosenthal et al. 2015; Hayward et al. 2011). One further model named c6e was a massive gas-rich galaxy with a halo mass of  $9 \times 10^{12} M_\odot$  and a gas fraction of 60 per cent, meant to mimic some submillimeter galaxy (SMG) properties. Fig. 7 shows the simulated SED for the M3–M3 merger case. We created output files at the specified intervals during the mergers of all combinations of the six galaxy models and then ran SUNRISE to compute the SED for each step and seven viewing angles of every merger.



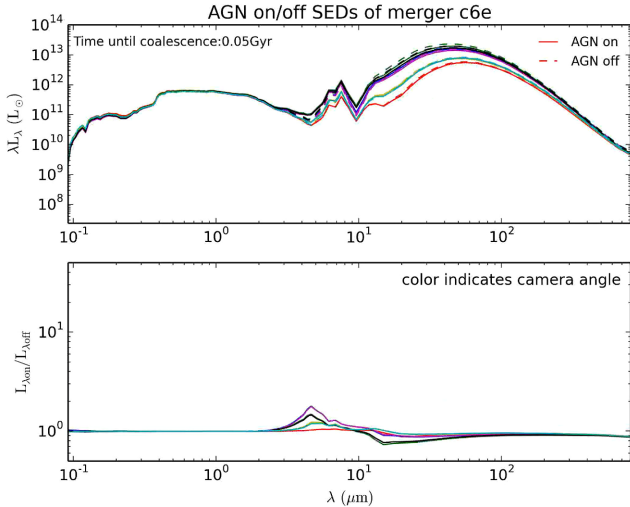
**Figure 8.** The SUNRISE SED output of the c6e–c6e merger. In this gas-rich example, the black hole separation is shown as the blue line, the star formation rate is indicated in red, and the black hole accretion rate is in green. The AGN fractional contribution to the luminosity is shown with the dotted black line; the  $L_{\text{AGN}}$  was calculated from the accretion rate, and  $L_{\text{SFR}}$  was calculated from the model’s SFR.

## 4.2 Simulated SEDs results

We did not see any evidence that CIGALE’s output reliability depended on the particulars of the merger scenario, and the M3–M3 or M4–M4 major merger simulations represent the results. As Lanz et al. (2014) have already described, those simulations give realistic SEDs (Weiner et al. in preparation). Fig. 7 (lower) illustrates the black hole accretion varying during the M3–M3 merger, peaking (for this example) at about  $0.7 M_\odot \text{yr}^{-1}$  shortly after the moment of coalescence. The accretion rate hovers at a few times  $10^{-3} M_\odot \text{yr}^{-1}$  for most of the early stages of the interaction, even during the first close pass of the two black hole nuclei, but starts to climb to its peak about 12 Myr before coalescence, when the separation is about 150 kpc. The accretion activity remains above the earlier baseline level for about 30 Myr, during which time the increasingly dense gas in the simulation also produces a burst of star formation, and afterwards the AGN accretion drops to a new baseline nearly 20 times smaller than the pre-merger level.

The gas-rich merger simulation ‘c6e–c6e’ has an initial gas fraction of 60 per cent, in contrast to the other simulated galaxies which, independent of mass, used gas fractions of only 15 and 38 per cent (Cox et al. 2008; Hayward et al. 2013; Lanz et al. 2014). The c6e–c6e simulation uses a baryonic mass of  $4 \times 10^{11} M_\odot$ , considerably more than the other simulations, but the same black hole mass of  $1.4 \times 10^5 M_\odot$ . For this merger, Fig. 8 plots the AGN fraction along with some other parameters versus time. The AGN luminosity in this example peaks briefly as high as 55 per cent at coalescence, and estimates of the SFR based solely on  $L_{\text{FIR}}$  will be correspondingly too high.

The SED simulations for c6e–c6e made with the AGN ‘turned off’, that is with  $\epsilon_r = 0$ , are illustrated in Fig. 9. At the largest viewing angle, (edge-on), the strongest difference between the two cases is a factor of two in the 5–8  $\mu\text{m}$  range. The fact that this part of the spectral range is most sensitive to the AGN fraction confirms what is already well-known from earlier *Spitzer* observations: the IRAC colour–colour diagram as manifest in the so-called Stern wedge is a useful tool to identify AGN (Stern et al. 2005). However, the reverse is not necessarily true, as Fig. 6 shows. Low-luminosity or highly obscured AGN may have blue [3.6]–[4.5] colours or red [5.8]–[8.0]



**Figure 9.** AGN on versus AGN off: the SUNRISE output for the c6e–c6e major merger at 0.7 Gyr. Colors correspond to different viewing angles. (Top) The SED for the AGN turned on (solid curves) and AGN turned off (dashed curves). (Bottom) The flux ratios for AGN on / AGN off, showing the spectral differences.

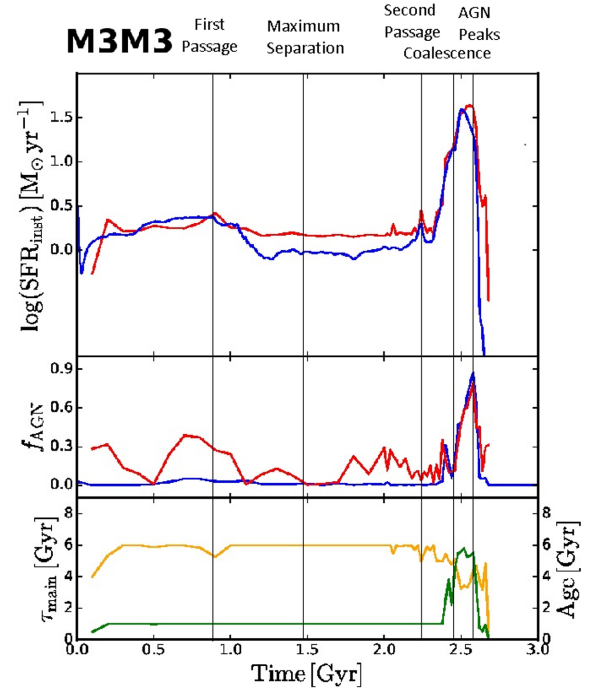
colours, and by using the full SED analysis we can obtain a more reliable accounting of AGN emission and demographics than IRAC colours alone.

#### 4.3 CIGALE performance

We compared CIGALE’s model results against simulations both with the AGN turned on (a normal  $\epsilon_r = 10$  per cent) and with the AGN turned off ( $\epsilon_r = 0$ ). The parameters of particular relevance here are: the ratio of the AGN’s dust torus radii, the optical depth at  $9.7 \mu\text{m}$ , the AGN opening angle, the AGN luminosity, and the AGN fraction as estimated both with the Dale et al. (2014) and the Fritz et al. (2006) methods. Fig. 10 plots the CIGALE-modeled outputs for the SFR and AGN fraction versus the simulated output values as a function of elapsed time with the AGN on.

CIGALE does an excellent job of evaluating  $f_{\text{AGN}}$  when the AGN is the dominant fraction of luminosity, but it does a poorer job of estimating  $f_{\text{AGN}}$  when it is below about 15 per cent. When CIGALE errs most often it overestimates  $f_{\text{AGN}}$ , but sometimes when  $f_{\text{AGN}} < 0.05$  CIGALE returns a value of zero. The CIGALE  $f_{\text{AGN}}$  estimate is particularly egregious in the 0.2–1.2 Gyr period, estimating a fraction of as much as 35 per cent when it is in fact less than 10 per cent. This occurs because slight changes to the relative flux densities in the 5–30  $\mu\text{m}$  bands happen in merger phases when the SFR is low and the longer-wavelength emission is therefore also weak. As a result, these slight changes have an undue impact on CIGALE by causing large shifts in the AGN fraction upward and correspondingly the SFR downward.

In low flux density cases, the full SED spectral output information is used to supplement the calculation of AGN luminosity fraction from the IR bands, which partially corrects for the  $f_{\text{AGN}}$  estimated by CIGALE. Similarly, when CIGALE underestimates the observed flux density at wavelengths below 10  $\mu\text{m}$  but overestimates it longward of 10 and 100  $\mu\text{m}$  (as we can see in the SED fit to UGC 5101, for example, in Fig. A1), additional correction can be applied by shifting upward the allowed range for the  $f_{\text{AGN}}$  parameter accordingly and/or using a cooler dust model. Not least, spectroscopic information (as per Sec. refssec:spec) can also be used



**Figure 10.** CIGALE-derived parameters for the M3–M3 merger; the top panel shows SFR, the middle panel shows  $f_{\text{AGN}}$ , and the bottom panel shows  $\tau_{\text{main}}$  and age. The blue curves show the model values output from the GADGET-3 simulations, and the red curves show values derived by CIGALE. The yellow line in the bottom panel is the derived age, and the green is the derived  $e$ -folding time for the delayed star formation history.

to calibrate a CIGALE output when the AGN fraction is low. In all of the AGN-dominated systems considered here, the AGN fraction is high enough that the CIGALE results are credible.

## 5 DISCUSSION

As shown in Fig. 3, the luminosity is strongly correlated with merger stage, which is expected. The AGN luminosity fraction is also correlated with both, as all of the stage 3.5 or less merging galaxies except for NGC 5394 have AGN luminosity fractions below the values of all of the stage 4 or higher merging galaxies. The lowest  $f_{\text{AGN}}$  values of 0 per cent were found in early-stage merging galaxies, and the late-stage merging galaxy with the lowest AGN fraction was Mrk 231, which is classified as a stage 6 at post-coalescence. This galaxy has already completed the merging process, and the AGN luminosity is dropping while the SFR is still relatively high.

Nearly half of the late-stage merging galaxies in our sample of 11 host an AGN that is at least as luminous as the rest of the galaxy in the IR. These behemoths can even exceed the IR luminosities of their hosts by an order of magnitude, as in IRAS 08572+3915, which at an estimated 91 per cent surpasses the original limit we had imposed on the AGN fraction at 90 per cent. This is consistent with Efstathiou et al. (2014), who have also reported an AGN fraction of  $\sim 90$  per cent. More comments on each galaxy in both samples can be found in Appendix A.

Our results are consistent with earlier analysis of merging galaxies. Timelines have been created of the luminosity of a model merging galaxy, highlighting the period of time when the AGN turns on in the galaxy’s later stages, near coalescence (e.g. Lanz et al. 2014; Hayward et al. 2014a). Our results are consistent with the timeline,



as the late-stage mergers are more likely to contain an AGN than the early-stage mergers in our sample, and the AGN fractions in the late-stage mergers are usually higher than those in the early-stage mergers.

Fig. 4 shows that the line ratios in Table 3 for the early-stage mergers do not correlate with the AGN fraction, but they do show a slight linear trend for the late-stage mergers. Overall, there was no measurable significant correlation between the  $[\text{Ne V}]/[\text{Ne II}]$  or  $[\text{O IV}]/[\text{Ne II}]$  ratio and the AGN fraction in the combined sample. This could be due to limitations in the measurements, as there were many upper limits for the line ratios in the early-stage merging galaxies in the Reference Sample at low AGN fraction. Some line ratios in our sample are similar to the line ratios for galaxies shown by Dale et al. (2009), but their sample is low-luminosity and unlikely to contain strong AGNs.

## 6 CONCLUSIONS

The AGN fraction in our Late-Stage Sample is systematically and significantly higher than that measured in our Reference Sample. For the Late-Stage Sample,  $f_{\text{AGN}}$  ranges from 17 to 91 per cent. The 91 per cent estimate is for IRAS 08572+3915, a late-stage ULIRG. The 17 per cent estimate is for Mrk 231, which is classified as a stage 6 post-coalescence merger having a high SFR. In contrast, in the Reference Sample,  $f_{\text{AGN}}$  is below 15 per cent for all but three galaxies. The difference is probably because the Late-Stage Sample galaxies are advanced in their interaction level, with material flowing to their centers and feeding the AGN that reside there, with the exception of MRK 231 which is consistent with being in the post-merger stage.

CIGALE SED modeling of late-stage snapshots of a set of SPH merger simulations yields AGN luminosity fractions that are in good agreement in general with the simulation values and also are consistent with values measured in the Late-Stage Merger Sample. However, CIGALE incorrectly inferred AGN fractions up to 30 per cent in earlier stages of the simulated mergers when the true value was close to 0. Galaxies in the Reference Sample with little to no empirical evidence in the SED for AGN activity were not modeled with large  $f_{\text{AGN}}$ , so the SED fits for these galaxies are reliable.

We also measured spectral line ratios for  $[\text{Ne V}]/[\text{Ne II}]$  and  $[\text{O IV}]/[\text{Ne II}]$  to provide another tool to estimate the strength of the AGN. We found no overall correlation above the  $2\sigma$  level in our samples; some strong AGNs have comparatively weak line ratios, similar to those of the weaker AGN. The effects of extinction in these mid-IR lines likely play a significant role. We do, however, find that the late-stage merging galaxies alone do show a possible linear trend between AGN fraction and line ratios.

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## APPENDIX A: NOTES ON INDIVIDUAL GALAXIES

### Group I: The Late-Stage Merger Sample

**IRAS 08572+3915:** IRAS 08572+3915 is a ULIRG. Its very steep spectrum from 2 to 20  $\mu\text{m}$  implies an extremely powerful AGN, which approaches 91 per cent of the total IR luminosity coming from the galaxy. This is consistent with Efstathiou et al. (2014), who also found an AGN luminosity fraction around 0.9, and Dale et al. (2014), who estimated an AGN contribution of 85 per cent.

**IRAS 15250+3609:** IRAS 15250+3609 is a ULIRG. The 2–10  $\mu\text{m}$  slope is steep enough to imply an AGN contribution to the SED, which CIGALE estimates at  $\sim 47$  per cent of the IR luminosity. Franceschini et al. (2003) define this galaxy as not AGN-dominated but still containing a LINER-type nucleus.

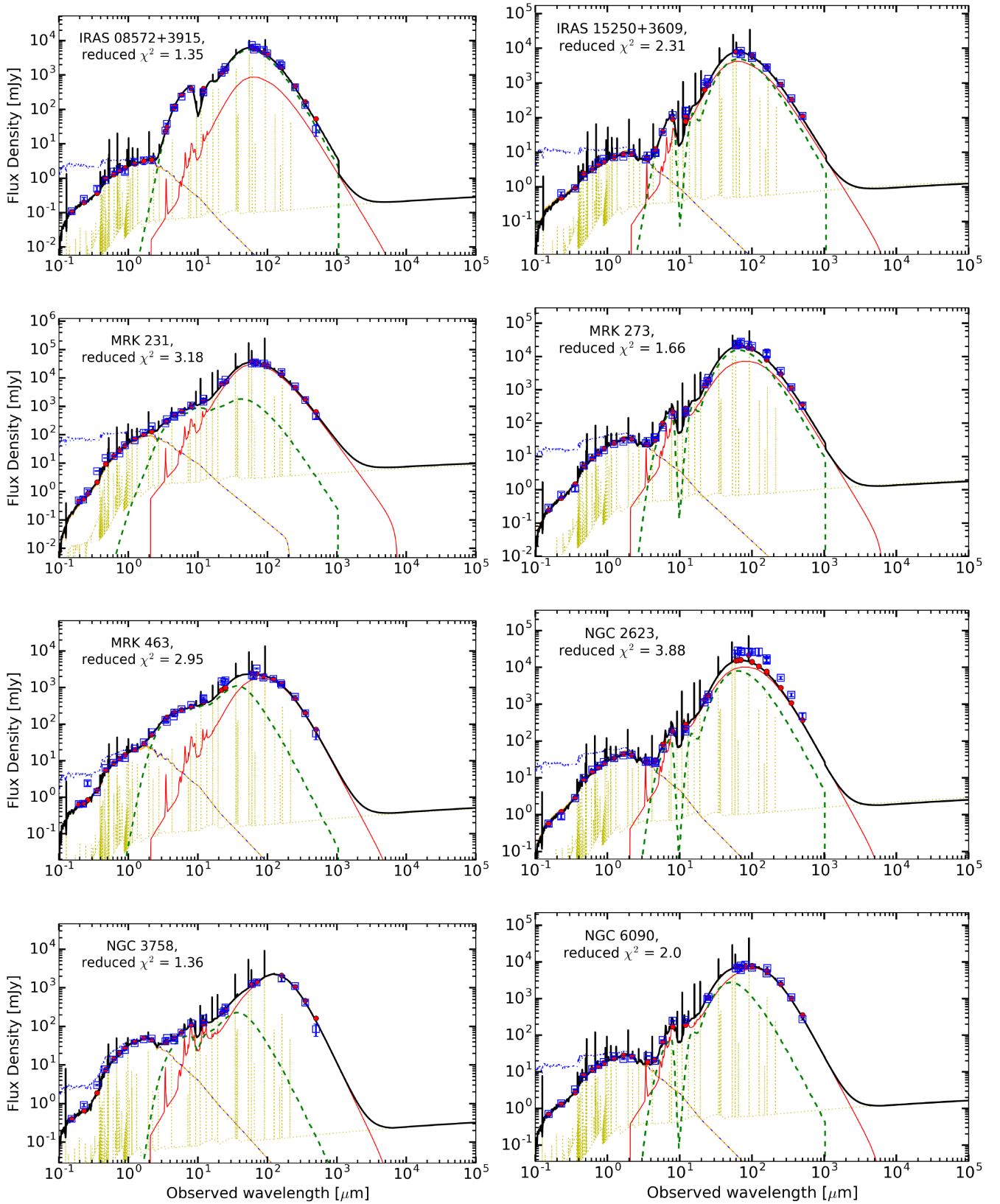
**Mrk 231:** Mrk 231 is a ULIRG that is the most IR-luminous system in the sample. We measure the AGN contribution at  $\sim 17$  per cent of the IR luminosity, the lowest value in the late-stage sample, lower than the value found by Rosenberg et al. (2015) by almost a factor of 5, and lower than that from Fritz et al. (2006) by almost a factor of 2. This low fraction is because the data at wavelengths greater than 60  $\mu\text{m}$  are well fit by the dust model with no AGN contribution needed.

**Mrk 273:** Mrk 273 is a ULIRG. The steep 3–24  $\mu\text{m}$  spectrum implies a large AGN contribution that CIGALE estimates at  $\sim 66$  per cent of the IR luminosity. This is higher than the value given by Rosenberg et al. (2015) by around a factor of 2.

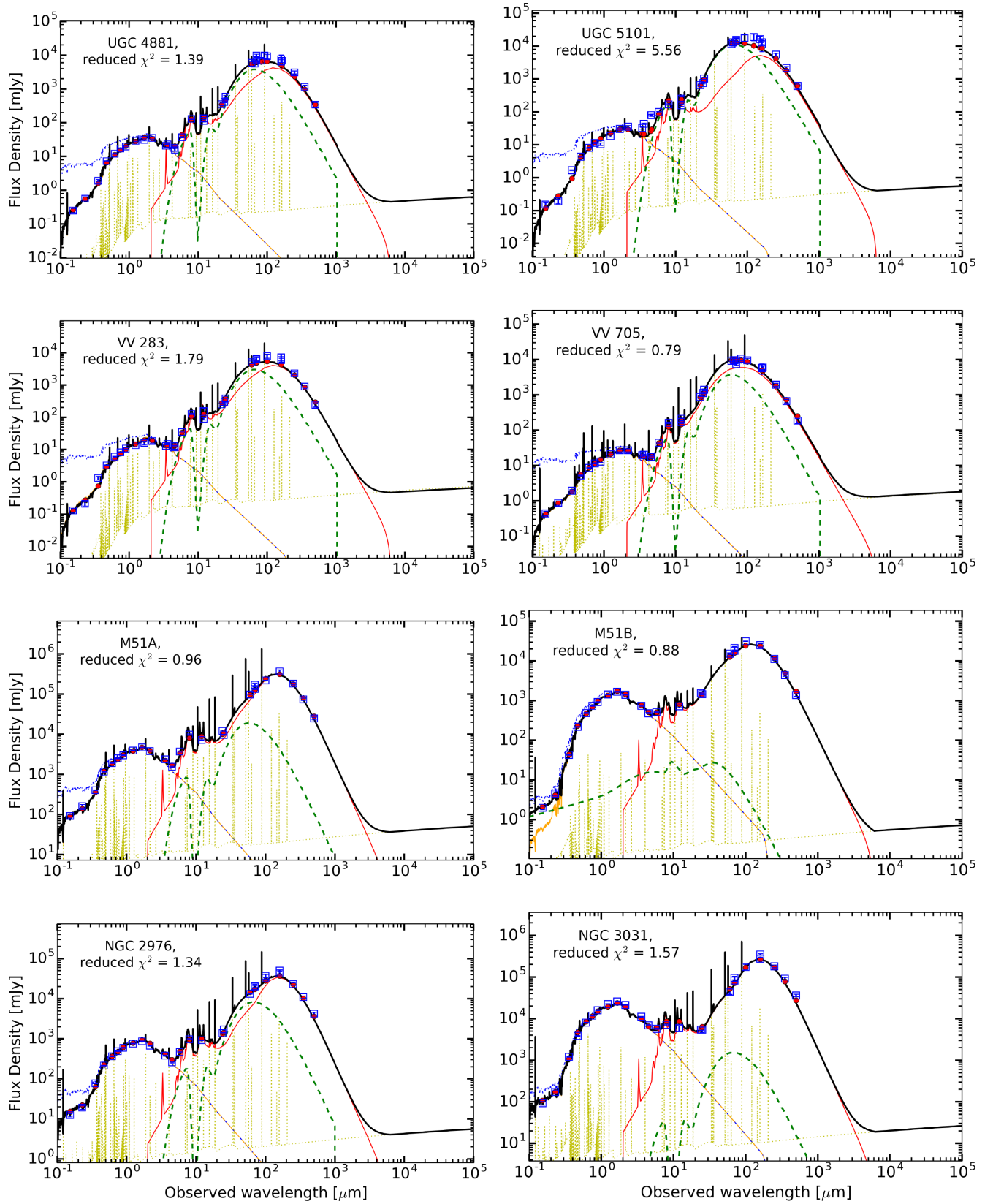
**Mrk 463:** Mrk 463 is an LIRG. The 3–24  $\mu\text{m}$  SED is well fit by an AGN model, and the derived AGN contribution is  $\sim 68$  per cent of the IR luminosity. The spectral line ratios for [Ne v]/[Ne ii] and [O iv]/[Ne ii] are a factor of 5 larger than others in this sample due to relatively low [Ne ii] flux, indicating that very little star formation is occurring. That the line ratio is so particularly strong is surprising, because an AGN strong enough to ionize neon and oxygen that heavily would also be expected to have powerful UV emission, which is not seen in Mrk 463. That no enhanced UV emission is seen is presumably due to a high internal extinction. Mrk 463 is also a luminous X-ray source, another indication of strong AGN activity.

**NGC 2623:** NGC 2623 is an LIRG with the 3.6–24  $\mu\text{m}$  data well fit by an AGN model. We measure the AGN contribution to be  $\sim 39$  per cent, but the  $\chi^2$  value is relatively large.

**NGC 3758:** NGC 3758 is an LIRG and the least luminous galaxy in the late-stage merger sample. The 3.6–24  $\mu\text{m}$  SED is well fit by an AGN model, but the estimated AGN contribution is  $\sim 30$  per cent.



**Figure A1.** Best-fitting SED models for the 24 galaxies in the sample containing the nebular emission (gold dotted lines), both attenuated stellar emission (orange) and non-attenuated stellar emission (blue dot-dashed), dust emission (red solid), and AGN emission (green dashed). The red dots are the best model flux densities and the blue squares mark the observed flux densities with  $1\sigma$  error bars.

Figure A1 – *continued*

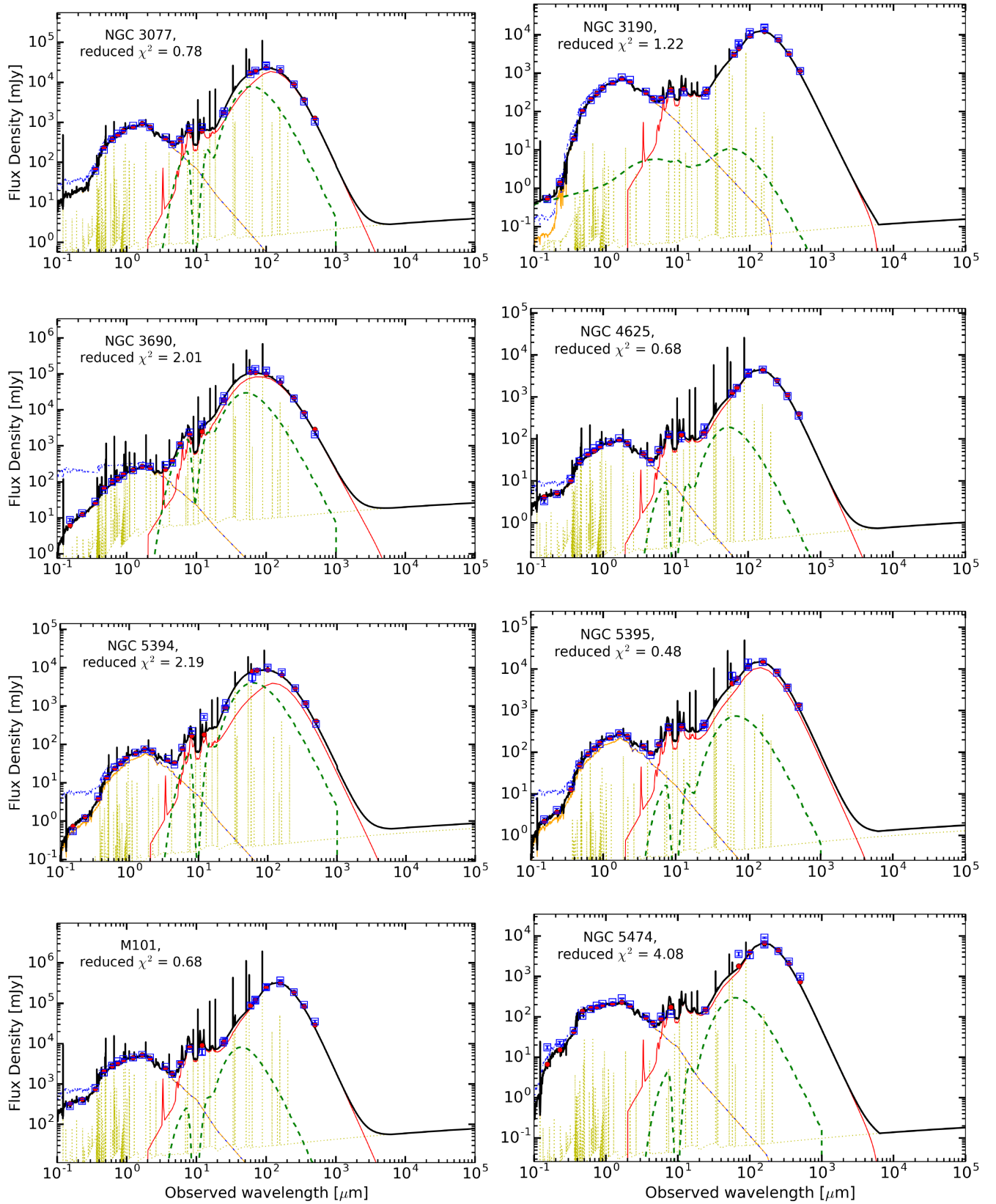


Figure A1 – continued



**NGC 6090:** NGC 6090 is an LIRG. The steep 4.5–8.0  $\mu\text{m}$  SED implies an AGN is present, but the estimated AGN luminosity fraction is  $\sim 26$  per cent, one of the lowest in the late-stage merger sample. This is still higher than the 10 per cent value measured by Dale et al. (2014).

**UGC 4881:** UGC 4881 is an LIRG. We measure the AGN contribution to be  $\sim 51$  per cent, although Dale et al. (2014) modeled it without an AGN. The red 5.8–100  $\mu\text{m}$  colours suggest that an AGN is present.

**UGC 5101:** UGC 5101 is a ULIRG. We measure the AGN contribution to be  $\sim 76$  per cent, although the best-fitting model is markedly worse than most of the late-stage mergers. Dale et al. (2014) estimated a value nearly a factor of 5 smaller. The red 4.5–24  $\mu\text{m}$  SED implies that an AGN is present.

**VV 283:** VV 283 is an LIRG. We measure an AGN contribution of  $\sim 47$  per cent, and the red 5.8–24  $\mu\text{m}$  SED implies that an AGN is present.

**VV 705:** VV 705 is an LIRG. We measure an AGN contribution of  $\sim 25$  per cent. This is corroborated by the results of Dale et al. (2014), who measured an AGN luminosity fraction of 25 per cent. The red 5.8–24  $\mu\text{m}$  SED implies that an AGN is present.

### Group II: The Reference Sample

**M51A:** M51A, also known as the Whirlpool Galaxy, is a well-known spiral galaxy with an elliptical companion. CIGALE modeled M51A with an AGN at 9 per cent of the IR luminosity. Hernández-García et al. (2016) modeled it as an obscured AGN, while L13 calculated a best-fitting model that did not have an AGN. Nothing in the SED demands the presence of an AGN.

**M51B:** M51B is the companion to M51A, and it was modeled with an AGN at  $< 3$  per cent of the IR luminosity. Hernández-García et al. (2016) classified it as a LINER galaxy, but L13 did not calculate an AGN luminosity fraction for it, and nothing in the SED requires an AGN to be present.

**NGC 2976:** NGC 2976 is a spiral galaxy in the M81 group. CIGALE modeled NGC 2976 with  $f_{\text{AGN}} = 28$  per cent, but previous results from González-Martín et al. (2015) and L13 produced models without AGNs. Nothing in the SED requires an AGN.

**NGC 3031:** NGC 3031, also known as M81, is a nearby spiral galaxy. CIGALE modeled the galaxy with  $f_{\text{AGN}} \leq 0.01$ , although L13 reported a total IR AGN luminosity fraction of 4 per cent and a maximum of 16 per cent in the 8–35  $\mu\text{m}$  range. The galaxy nucleus has a unique dust spectrum (Smith et al. 2010), and modeling based on standard templates is unreliable. However, nothing in the total-galaxy SED requires an AGN.

**NGC 3077:** NGC 3077 is a low-luminosity irregular galaxy. We measure the AGN contribution at  $\sim 34$  per cent, which is not consistent with results from Hernández-García et al. (2016) and L13. NGC 3077 is a fairly isolated companion of NGC 3031 = M81, 46' away (Knapen et al. 2014), indicating it is in the earliest

stages of merging. It does show signs of previous galaxy interaction, but nothing in the SED requires an AGN.

**NGC 3190:** NGC 3190 is an edge-on spiral galaxy with prominent dust lanes. CIGALE places a  $3\sigma$  upper limit for the AGN contribution at 1 per cent, although it has been shown to have a LINER-type nucleus (González-Martín et al. 2015). L13 also described a best-fitting model with no AGN contribution, and nothing in the SED suggests an AGN.

**NGC 3690:** NGC 3690 is the most IR-luminous galaxy in the Reference Sample and its only LIRG (Rosenberg et al. 2014). This galaxy is nearing final pass; although not at coalescence, it is still classified as a late-stage merger. We measure the AGN contribution at  $\sim 30$  per cent. Vardoulaki et al. (2015) have shown it to be a composite galaxy, containing a LINER-type nucleus while also undergoing SF. L13 described the SF but did not mention an AGN. Dale et al. (2014) characterized the AGN contribution to the SED at 50 per cent. The red 3.6–24  $\mu\text{m}$  SED suggests an AGN is present.

**NGC 4625:** NGC 4625 is a peculiar spiral with a blue SED. CIGALE fits  $f_{\text{AGN}} \leq 3$  per cent. Véron-Cetty & Véron (2010) classified it as a Seyfert galaxy (type unknown), but L13 did not define an AGN contribution for NGC 4625. Nothing in the SED requires an AGN.

**NGC 5394:** NGC 5394 is a companion of NGC 5395 in the middle stages of merging at a projected separation of 28 kpc. We measure the AGN contribution at  $\sim 58$  per cent. Toba et al. (2013) stated NGC 5394 as a composite, although L13 did not place an AGN in NGC 5394. The red 4.5–24  $\mu\text{m}$  SED suggests an AGN is present.

**NGC 5395:** NGC 5395 is the larger spiral companion of NGC 5394. CIGALE fits  $f_{\text{AGN}} < 10$  per cent to NGC 5395, consistent with both Véron-Cetty & Véron (2010) and L13 calling it a LINER and attributing to it 3–12 per cent of the bolometric and mid-IR luminosities. Nothing in the SED requires an AGN.

**M101:** M101, also known as the Pinwheel Galaxy, is a nearby spiral galaxy showing some tidal disruptions in its outer arms with multiple small companions, including NGC 5474. CIGALE posits  $f_{\text{AGN}} < 14$  per cent, although Brassington et al. (2015) defined it as not having an AGN, and L13 also gave a best-fitting model with no AGN. The red 5.8–8.0 colour is consistent with either an AGN or high SFR.

**NGC 5474:** NGC 5474 is a smaller companion to M101 at a projected separation of 87 kpc. CIGALE models NGC 5474 with  $f_{\text{AGN}} \leq 5$  per cent. Brassington et al. (2015) and L13 did not fit a model containing an AGN to the data, and nothing in the SED requires an AGN.

## APPENDIX B: PHOTOMETRY AND PACS SPECTROPHOTOMETRIC DATA

**Table B1.** *GALEX* and *Swift* UV photometry.

Galaxy Name	<i>GALEX</i>		<i>Swift</i>		
	FUV (mJy)	NUV (mJy)	UVOT_UVW2 (mJy)	UVOT_UVM2 (mJy)	UVOT_UVW1 (mJy)
Late-Stage Merger Sample					
IRAS 08572+3915	0.110 ± 0.012	0.130 ± 0.014	...	...	...
IRAS 15250+3609	...	0.652 ± 0.066	...	...	...
Mrk 231	...	...	0.487 ± 0.050	0.531 ± 0.056	1.02 ± 0.11
Mrk 273	0.290 ± 0.041	0.71 ± 0.11	...	...	...
Mrk 463	...	...	0.645 ± 0.066	0.659 ± 0.069	2.42 ± 0.24
NGC 2623	0.58 ± 0.06	0.89 ± 0.09	...	...	...
NGC 3758	0.403 ± 0.040	0.919 ± 0.092	...	...	...
NGC 6090	0.705 ± 0.071	1.37 ± 0.14	...	...	...
UGC 4881	0.250 ± 0.036	0.56 ± 0.08	...	...	...
UGC 5101	0.150 ± 0.021	0.190 ± 0.031	...	...	...
VV 283	0.130 ± 0.019	0.210 ± 0.032	...	...	...
VV 705	0.440 ± 0.062	0.87 ± 0.13	...	...	...
Reference Sample					
M51A	90.9 ± 9.09	162.0 ± 16.2	...	1110. ± 63	1650. ± 95
M51B	1.98 ± 0.20	4.53 ± 0.45	89.90 ± 6.23	282.0 ± 18.7	551.0 ± 37.1
NGC 2976	12.4 ± 1.24	19.2 ± 1.92	85.10 ± 10.90	200.0 ± 25.5	314.0 ± 40.1
NGC 3031	93.0 ± 9.0	173.0 ± 17.0	...	2970. ± 83	6100. ± 180.
NGC 3077	...	...	90.70 ± 11.70	243.0 ± 30.9	418.0 ± 53.3
NGC 3190	0.53 ± 0.05	1.46 ± 0.15	16.50 ± 1.63	60.50 ± 5.83	126.0 ± 12.2
NGC 3690	8.76 ± 0.88	13.6 ± 1.36	...	...	...
NGC 4625	3.27 ± 0.33	4.88 ± 0.49	14.50 ± 0.88	28.90 ± 1.09	41.80 ± 1.76
NGC 5394	0.55 ± 0.05	1.20 ± 0.12	5.37 ± 1.16	14.10 ± 2.85	22.80 ± 4.66
NGC 5395	2.42 ± 0.24	4.05 ± 0.41	23.90 ± 5.06	61.60 ± 12.50	103.0 ± 21.1
M101	283.0 ± 28.0	380.0 ± 38.0	...	2020. ± 175	2610. ± 248
NGC 5474	17.5 ± 1.75	22.5 ± 2.25	...	131.0 ± 19.4	176.0 ± 27.6

**Table B2.** SDSS photometry.

Galaxy name	SDSS				
	<i>u</i> (mJy)	<i>g</i> (mJy)	<i>r</i> (mJy)	<i>i</i> (mJy)	<i>z</i> (mJy)
Late-Stage Merger Sample					
IRAS 08572+3915	0.510 ± 0.075	0.870 ± 0.028	1.58 ± 0.047	1.61 ± 0.068	1.59 ± 0.20
IRAS 15250+3609	0.906 ± 0.084	1.96 ± 0.055	3.25 ± 0.088	...	...
Mrk 231	5.34 ± 0.23	14.90 ± 0.43	17.10 ± 0.48	20.90 ± 0.59	39.20 ± 1.29
Mrk 273	1.09 ± 0.19	5.29 ± 0.17	9.23 ± 0.28	13.10 ± 0.41	15.50 ± 0.83
Mrk 463	1.39 ± 0.12	6.27 ± 0.13	8.81 ± 0.19	12.00 ± 0.26	11.80 ± 0.50
NGC 2623	2.84 ± 0.15	9.95 ± 0.20	16.69 ± 0.34	21.3 ± 0.44	26.8 ± 0.78
NGC 3758	3.02 ± 0.13	7.72 ± 0.16	13.80 ± 0.29	19.60 ± 0.41	24.50 ± 0.61
NGC 6090	3.31 ± 0.08	7.37 ± 0.15	11.90 ± 0.24	15.50 ± 0.32	18.60 ± 0.47
UGC 4881	1.92 ± 0.16	7.23 ± 0.21	12.10 ± 0.36	17.30 ± 0.52	19.80 ± 0.87
UGC 5101	1.680 ± 0.051	4.35 ± 0.15	8.06 ± 0.26	11.30 ± 0.38	10.4 ± 1.0
VV 283	1.32 ± 0.17	3.25 ± 0.11	6.07 ± 0.19	8.25 ± 0.27	10.80 ± 0.54
VV 705	1.92 ± 0.14	5.18 ± 0.16	8.55 ± 0.26	11.50 ± 0.35	12.90 ± 0.60
Reference Sample					
M51A	402.0 ± 8.1	1170 ± 23	1920 ± 38	2520 ± 50	2980 ± 60
M51B	42.7 ± 0.88	218.0 ± 4.4	484.0 ± 9.7	734.0 ± 14.7	920.0 ± 18.4
NGC 2976	77.1 ± 1.57	213.0 ± 4.3	366.0 ± 7.3	488.0 ± 9.8	593.0 ± 11.9
NGC 3031	1210 ± 24	3850 ± 77	7980 ± 160	11800 ± 240	15900 ± 320
NGC 3077	61.4 ± 1.3	205.0 ± 4.0	381.0 ± 8.0	528.0 ± 11.0	617.0 ± 12.0
NGC 3190	21.0 ± 0.44	94.7 ± 1.9	198.0 ± 4.0	299.0 ± 6.0	394.0 ± 7.9
NGC 3690	28.5 ± 0.59	58.7 ± 1.2	100.0 ± 2.0	120.0 ± 2.4	148.0 ± 3.0
NGC 4625	11.3 ± 0.26	29.4 ± 0.59	47.1 ± 0.94	60.1 ± 1.21	71.1 ± 1.48
NGC 5394	4.38 ± 0.13	14.6 ± 0.29	25.1 ± 0.51	34.0 ± 0.69	39.9 ± 0.92
NGC 5395	15.0 ± 0.33	51.0 ± 1.02	93.0 ± 1.86	130.0 ± 2.6	159.0 ± 3.2
M101	800.0 ± 16.0	1920 ± 38	2860 ± 57	3650 ± 73	4160 ± 83
NGC 5474	44.9 ± 0.93	106.0 ± 2.0	155.0 ± 3.0	187.0 ± 4.0	209.0 ± 4.0

**Table B3.** 2MASS and IRAS photometry.

Galaxy name	2MASS			IRAS			
	<i>J</i> (mJy)	<i>H</i> (mJy)	<i>K<sub>s</sub></i> (mJy)	12 $\mu$ m (mJy)	25 $\mu$ m (mJy)	60 $\mu$ m (mJy)	100 $\mu$ m (mJy)
Late-Stage Merger Sample							
IRAS 08572+3915	2.91 $\pm$ 0.51	3.91 $\pm$ 0.84	3.87 $\pm$ 0.82	318 $\pm$ 35	1700 $\pm$ 90	7430 $\pm$ 370	4770 $\pm$ 150
IRAS 15250+3609	7.28 $\pm$ 0.15	7.46 $\pm$ 0.15	9.56 $\pm$ 0.19	160 $\pm$ 30	1310 $\pm$ 25	7200 $\pm$ 40	5930 $\pm$ 98
Mrk 231	60.8 $\pm$ 2.2	111.0 $\pm$ 3.8	192.0 $\pm$ 6.1	1870 $\pm$ 90	8660 $\pm$ 430	32000 $\pm$ 1600	29700 $\pm$ 1000
Mrk 273	23.7 $\pm$ 1.7	27.9 $\pm$ 2.5	32.6 $\pm$ 2.8	235 $\pm$ 27	2280 $\pm$ 130	21700 $\pm$ 870	22500 $\pm$ 900
Mrk 463	21.5 $\pm$ 1.2	31.8 $\pm$ 2.0	60.2 $\pm$ 2.1	510 $\pm$ 40	1580 $\pm$ 90	2180 $\pm$ 110	1920 $\pm$ 210
NGC 2623	34.3 $\pm$ 1.2	40.8 $\pm$ 1.7	42.6 $\pm$ 1.8	210 $\pm$ 20	1810 $\pm$ 40	23700 $\pm$ 930	25900 $\pm$ 1000
NGC 3758	36.8 $\pm$ 1.3	46.4 $\pm$ 1.8	48.4 $\pm$ 2.0	160 $\pm$ 30	309 $\pm$ 43	1260 $\pm$ 130	2410 $\pm$ 190
NGC 6090	26.5 $\pm$ 0.8	22.5 $\pm$ 1.0	23.5 $\pm$ 1.2	260 $\pm$ 20	1110 $\pm$ 40	6660 $\pm$ 270	9400 $\pm$ 1000
UGC 4881	29.1 $\pm$ 1.7	31.2 $\pm$ 2.2	35.8 $\pm$ 2.8	135 $\pm$ 31	599 $\pm$ 48	5960 $\pm$ 360	10300 $\pm$ 1100
UGC 5101	18.6 $\pm$ 1.7	30.1 $\pm$ 2.5	35.3 $\pm$ 2.9	250 $\pm$ 40	1030 $\pm$ 60	11500 $\pm$ 810	19900 $\pm$ 1400
VV 283	13.7 $\pm$ 1.5	16.0 $\pm$ 2.1	18.5 $\pm$ 2.8	157 $\pm$ 33	386 $\pm$ 66	5070 $\pm$ 460	7950 $\pm$ 480
VV 705	20.8 $\pm$ 1.6	26.9 $\pm$ 2.5	26.9 $\pm$ 2.7	210 $\pm$ 20	1390 $\pm$ 70	9210 $\pm$ 370	10000 $\pm$ 900
Reference Sample							
M51A	3940 $\pm$ 79	4690 $\pm$ 94	3810 $\pm$ 77	7210 $\pm$ 75	9560 $\pm$ 77	97400 $\pm$ 190	221000 $\pm$ 300
M51B	1420 $\pm$ 28	1670 $\pm$ 34	1400 $\pm$ 28	721.0 $\pm$ 57.7	1450 $\pm$ 51	15200 $\pm$ 800	31300 $\pm$ 370
NGC 2976	757.0 $\pm$ 15.3	862.0 $\pm$ 17.6	670.0 $\pm$ 14.1	920.0 $\pm$ 20.0	1710 $\pm$ 20	13100 $\pm$ 30	33400 $\pm$ 340
NGC 3031	22300 $\pm$ 446	25700 $\pm$ 515	21300 $\pm$ 427	5860 $\pm$ 879	5420 $\pm$ 813	44700 $\pm$ 6710	174000 $\pm$ 26100
NGC 3077	826.0 $\pm$ 17.0	937.0 $\pm$ 19.0	759.0 $\pm$ 16.0	760.0 $\pm$ 23.0	1880 $\pm$ 25.0	15900 $\pm$ 390	26500 $\pm$ 1320
NGC 3190	602.0 $\pm$ 12.1	773.0 $\pm$ 15.6	661.0 $\pm$ 13.4	315.0 $\pm$ 31.5	351.0 $\pm$ 75.7	3190 $\pm$ 35	10100 $\pm$ 510
NGC 3690	212.0 $\pm$ 4.4	278.0 $\pm$ 5.9	262.0 $\pm$ 5.5	3900 $\pm$ 400	24100 $\pm$ 2400	122000 $\pm$ 12500	123000 $\pm$ 12500
NGC 4625	83.1 $\pm$ 1.97	104.0 $\pm$ 2.7	78.7 $\pm$ 2.4	117.0 $\pm$ 31.6	188.0 $\pm$ 24.6	1200 $\pm$ 132	3580 $\pm$ 250
NGC 5394	60.5 $\pm$ 1.36	74.4 $\pm$ 1.76	65.1 $\pm$ 1.75	520.0 $\pm$ 50.0	1190 $\pm$ 110	5620 $\pm$ 1410	...
NGC 5395	237.0 $\pm$ 4.8	285.0 $\pm$ 5.9	240.0 $\pm$ 5.1	400.0 $\pm$ 40.0	480.0 $\pm$ 60.0	6860 $\pm$ 1500	14200 $\pm$ 3100
M101	4540 $\pm$ 92	5270 $\pm$ 107	4570 $\pm$ 94	6200 $\pm$ 930	11800 $\pm$ 1770	88000 $\pm$ 13200	253000 $\pm$ 37900
NGC 5474	229.0 $\pm$ 5.0	288.0 $\pm$ 7.0	201.0 $\pm$ 6.0	...	...	1330 $\pm$ 67	...

**Table B4.** WISE photometry for the late-stage merger sample; it was not used for the reference sample (see Section 2).

Galaxy name	3.4 $\mu$ m (mJy)	4.6 $\mu$ m (mJy)	12 $\mu$ m (mJy)	22 $\mu$ m (mJy)
Late-Stage Merger Sample				
IRAS 08572+3915	24.0 $\pm$ 1.44	113.0 $\pm$ 6.78	308.0 $\pm$ 18.5	1170 $\pm$ 70
IRAS 15250+3609	6.29 $\pm$ 0.37	9.64 $\pm$ 0.37	122.0 $\pm$ 37.3	946.0 $\pm$ 56.7
Mrk 231	246.0 $\pm$ 20.9	420.0 $\pm$ 35.6	1360.0 $\pm$ 115.0	6080 $\pm$ 530
Mrk 273	24.8 $\pm$ 1.58	36.4 $\pm$ 2.45	212.0 $\pm$ 14.2	1460 $\pm$ 130
Mrk 463	131.0 $\pm$ 7.86	206.0 $\pm$ 12.4	476.0 $\pm$ 28.6	1490 $\pm$ 90
NGC 2623	27.4 $\pm$ 1.64	25.4 $\pm$ 1.52	181.3 $\pm$ 10.9	1120 $\pm$ 70
NGC 3758	40.9 $\pm$ 2.46	43.4 $\pm$ 2.61	127.0 $\pm$ 7.63	227.0 $\pm$ 14.2
NGC 6090	...	...	...	...
UGC 4881	20.7 $\pm$ 1.76	15.8 $\pm$ 1.35	114.0 $\pm$ 9.69	352.0 $\pm$ 30.7
UGC 5101	32.4 $\pm$ 2.75	79.0 $\pm$ 6.71	158.0 $\pm$ 13.4	646.0 $\pm$ 56.8
VV 283	13.0 $\pm$ 1.11	11.8 $\pm$ 1.51	89.6 $\pm$ 7.61	252.0 $\pm$ 22.1
VV 705	19.5 $\pm$ 1.65	17.2 $\pm$ 1.47	169.0 $\pm$ 14.4	873.0 $\pm$ 81.9

**Table B5.** *Spitzer*/IRAC and MIPS photometry.

Galaxy name	IRAC				MIPS		
	3.6 $\mu\text{m}$ (mJy)	4.5 $\mu\text{m}$ (mJy)	5.8 $\mu\text{m}$ (mJy)	8.0 $\mu\text{m}$ (mJy)	24 $\mu\text{m}$ (mJy)	70 $\mu\text{m}$ (mJy)	160 $\mu\text{m}$ (mJy)
Late-Stage Merger Sample							
IRAS 08572+3915	38.6 $\pm$ 1.2	98.3 $\pm$ 3.0	257.0 $\pm$ 7.7	350 $\pm$ 10	1390 $\pm$ 56	6160 $\pm$ 250	1850 $\pm$ 74
IRAS 15250+3609	8.02 $\pm$ 0.22	9.84 $\pm$ 0.29	38.3 $\pm$ 1.2	118.0 $\pm$ 3.5	1050 $\pm$ 40	8410 $\pm$ 330	2750 $\pm$ 110
Mrk 231	357.0 $\pm$ 10.9	473.0 $\pm$ 14.4	626.0 $\pm$ 19.7	1070 $\pm$ 30	...	...	...
Mrk 273	28.8 $\pm$ 1.9	37.9 $\pm$ 1.6	75.2 $\pm$ 5.1	177.0 $\pm$ 7.5	1800 $\pm$ 70	27400 $\pm$ 4700	12500 $\pm$ 4100
Mrk 463	116.0 $\pm$ 3.5	163.0 $\pm$ 4.9	256.0 $\pm$ 7.7	332.0 $\pm$ 10.0	1430 $\pm$ 60	3310 $\pm$ 140	1190 $\pm$ 50
NGC 2623	29.2 $\pm$ 0.88	27.4 $\pm$ 0.82	62.4 $\pm$ 1.9	178.0 $\pm$ 9.3	1390 $\pm$ 60	27300 $\pm$ 2200	15200 $\pm$ 1500
NGC 3758	50.3 $\pm$ 1.5	51.9 $\pm$ 1.6	69.1 $\pm$ 2.1	120.0 $\pm$ 3.6	250 $\pm$ 10	...	...
NGC 6090	27.0 $\pm$ 0.9	19.9 $\pm$ 0.6	73.6 $\pm$ 2.3	247.0 $\pm$ 7.4	970 $\pm$ 39.0	7890 $\pm$ 320	...
UGC 4881	22.9 $\pm$ 1.0	17.90 $\pm$ 0.76	37.0 $\pm$ 1.6	130.0 $\pm$ 5.5	415 $\pm$ 24	9740 $\pm$ 960	7600 $\pm$ 2500
UGC 5101	43.5 $\pm$ 1.9	79.2 $\pm$ 3.4	98.3 $\pm$ 4.2	168.0 $\pm$ 7.1	753 $\pm$ 43	15400 $\pm$ 1400	11300 $\pm$ 4500
VV 283	15.00 $\pm$ 0.64	13.00 $\pm$ 0.55	34.6 $\pm$ 1.5	128.0 $\pm$ 5.4	300 $\pm$ 17	6390 $\pm$ 730	7190 $\pm$ 2030
VV 705	20.00 $\pm$ 0.85	17.50 $\pm$ 0.75	43.5 $\pm$ 1.9	161.0 $\pm$ 6.8	1130 $\pm$ 64	7480 $\pm$ 650	5690 $\pm$ 1880
Reference Sample							
M51A	2370 $\pm$ 71	1550 $\pm$ 47	3690 $\pm$ 111	9810 $\pm$ 294	12000 $\pm$ 480	135000 $\pm$ 5400	...
M51B	739.0 $\pm$ 22.2	479.0 $\pm$ 14.4	515.0 $\pm$ 15.4	...	1540 $\pm$ 62	18000 $\pm$ 720	...
NGC 2976	378.0 $\pm$ 11.3	256.0 $\pm$ 7.7	465.0 $\pm$ 14.0	972.0 $\pm$ 29.2	1380 $\pm$ 55	20000 $\pm$ 800	50300 $\pm$ 2000
NGC 3031	11000 $\pm$ 330	6930 $\pm$ 208	5700 $\pm$ 171	7060 $\pm$ 212	5410 $\pm$ 216	82400 $\pm$ 3300	348000 $\pm$ 13900
NGC 3077	424.0 $\pm$ 13.0	284.0 $\pm$ 9.0	374.0 $\pm$ 11.0	716.0 $\pm$ 21.0	1650 $\pm$ 66	...	...
NGC 3190	330.0 $\pm$ 9.9	211.0 $\pm$ 6.3	206.0 $\pm$ 6.2	293.0 $\pm$ 8.8	262.0 $\pm$ 10.6	5530 $\pm$ 224	15400 $\pm$ 224
NGC 3690	295.0 $\pm$ 8.8	340.0 $\pm$ 10.2	1040 $\pm$ 31	2370 $\pm$ 71	17400 $\pm$ 700	...	...
NGC 4625	43.0 $\pm$ 2.6	28.1 $\pm$ 0.85	54.1 $\pm$ 1.6	126.0 $\pm$ 3.8	127.0 $\pm$ 5.2	...	...
NGC 5394	41.9 $\pm$ 1.26	29.6 $\pm$ 0.89	82.1 $\pm$ 2.47	222.0 $\pm$ 6.7	846.0 $\pm$ 33.9	...	...
NGC 5395	131.0 $\pm$ 3.9	85.9 $\pm$ 2.58	164.0 $\pm$ 4.9	396.0 $\pm$ 11.9	400.0 $\pm$ 16.1	...	...
M101	2660 $\pm$ 80	1770 $\pm$ 53	3110 $\pm$ 93	7470 $\pm$ 224	10500 $\pm$ 420	117000 $\pm$ 4700	369000 $\pm$ 14800
NGC 5474	101.0 $\pm$ 3.0	67.0 $\pm$ 2.0	101.0 $\pm$ 3.0	117.0 $\pm$ 4.0	143.0 $\pm$ 6.0	...	9050 $\pm$ 379



**Table B6.** *Herschel*/PACS and SPIRE photometry.

Galaxy name	PACS			SPIRE		
	75 $\mu\text{m}$ (mJy)	110 $\mu\text{m}$ (mJy)	170 $\mu\text{m}$ (mJy)	250 $\mu\text{m}$ (mJy)	350 $\mu\text{m}$ (mJy)	500 $\mu\text{m}$ (mJy)
Late-Stage Merger Sample						
IRAS 08572+3915	6190 $\pm$ 620	4120 $\pm$ 410	1830 $\pm$ 180	446 $\pm$ 33	131 $\pm$ 16	26.7 $\pm$ 10.3
IRAS 15250+3609	6720 $\pm$ 673	5230 $\pm$ 524	2780 $\pm$ 279	979 $\pm$ 73	359 $\pm$ 31	111 $\pm$ 15
Mrk 231	31400 $\pm$ 4500	26800 $\pm$ 3800	14900 $\pm$ 2100	4990 $\pm$ 500	1670 $\pm$ 180	456 $\pm$ 87
Mrk 273	23300 $\pm$ 2300	20600 $\pm$ 2200	11600 $\pm$ 1300	3710 $\pm$ 460	1200 $\pm$ 130	334 $\pm$ 52
Mrk 463	...	...	...	559 $\pm$ 45	199 $\pm$ 19	55.4 $\pm$ 18.1
NGC 2623	25900 $\pm$ 2600	26300 $\pm$ 2600	15800 $\pm$ 1600	5180 $\pm$ 0520	1790 $\pm$ 200	473 $\pm$ 90
NGC 3758	1390 $\pm$ 170	...	1740 $\pm$ 280	1100 $\pm$ 90	416 $\pm$ 42	85.1 $\pm$ 30.4
NGC 6090	6020 $\pm$ 600	7020 $\pm$ 700	4860 $\pm$ 490	2910 $\pm$ 200	1080 $\pm$ 80	280 $\pm$ 20
UGC 4881	7410 $\pm$ 1050	9270 $\pm$ 1310	6350 $\pm$ 930	3090 $\pm$ 320	1150 $\pm$ 140	344 $\pm$ 55
UGC 5101	14000 $\pm$ 2000	...	13000 $\pm$ 1900	5490 $\pm$ 550	2140 $\pm$ 220	610 $\pm$ 92
VV 283	6170 $\pm$ 880	7840 $\pm$ 1110	5900 $\pm$ 840	2320 $\pm$ 240	891 $\pm$ 97	244 $\pm$ 38
VV 705	9440 $\pm$ 1340	9420 $\pm$ 1330	5770 $\pm$ 820	1950 $\pm$ 200	667 $\pm$ 78	185 $\pm$ 29
Reference Sample						
M51A	170000 $\pm$ 17000	...	367000 $\pm$ 36700	183000 $\pm$ 12800	74300 $\pm$ 5200	24700 $\pm$ 1730
M51B	21100 $\pm$ 2110	...	24800 $\pm$ 2480	11200 $\pm$ 790	4360 $\pm$ 310	1380 $\pm$ 102
NGC 2976	20600 $\pm$ 2060	37000 $\pm$ 3700	42900 $\pm$ 4300	23000 $\pm$ 1610	10900 $\pm$ 800	4220 $\pm$ 300
NGC 3031	94700 $\pm$ 9500	...	272000 $\pm$ 27300	180000 $\pm$ 12600	91000 $\pm$ 6380	35900 $\pm$ 2520
NGC 3077	19500 $\pm$ 2000	25100 $\pm$ 2510	21200 $\pm$ 2130	8870 $\pm$ 626	3290 $\pm$ 241	1030 $\pm$ 84
NGC 3190	6020 $\pm$ 614	11800 $\pm$ 1180	14600 $\pm$ 1470	7950 $\pm$ 558	3420 $\pm$ 243	1130 $\pm$ 84
NGC 3690	138000 $\pm$ 13800	124000 $\pm$ 12400	69900 $\pm$ 6990	20800 $\pm$ 1460	7170 $\pm$ 503	2070 $\pm$ 146
NGC 4625	1640 $\pm$ 187	3810 $\pm$ 389	4450 $\pm$ 455	2240 $\pm$ 160	1050 $\pm$ 79	360.0 $\pm$ 31.8
NGC 5394	7860 $\pm$ 787	10000 $\pm$ 1010	7650 $\pm$ 766	2920 $\pm$ 205	1120 $\pm$ 81	339.0 $\pm$ 26.6
NGC 5395	5250 $\pm$ 528	11200 $\pm$ 1120	14800 $\pm$ 1480	8190 $\pm$ 574	3570 $\pm$ 251	1230 $\pm$ 88
M101	123000 $\pm$ 12300	249000 $\pm$ 24900	312000 $\pm$ 31200	189000 $\pm$ 13200	91400 $\pm$ 6400	35300 $\pm$ 2480
NGC 5474	3580 $\pm$ 403	3320 $\pm$ 582	6120 $\pm$ 648	4210 $\pm$ 304	2330 $\pm$ 176	985.0 $\pm$ 83.0

**Table B7.** The PACS spectrophotometric (SP) data points, with wavelength, number of observations, and average flux; ... signifies unknown number of observations were used.

Galaxy Name	PACS ‘Blue’ SP			PACS ‘Green’ SP			PACS ‘Red’ SP		
	$\lambda$ ( $\mu\text{m}$ )	#	Flux Density (mJy)	$\lambda$ ( $\mu\text{m}$ )	#	Flux Density (mJy)	$\lambda$ ( $\mu\text{m}$ )	#	Flux Density (mJy)
IRAS 08572+3915	64	386	5510 $\pm$ 550	89	497	5320 $\pm$ 530	151	60	2160 $\pm$ 220
Mrk 231	66	701	35700 $\pm$ 3570	90	714	32200 $\pm$ 3220	160	490	16200 $\pm$ 1620
Mrk 273	63	419	24400 $\pm$ 2440	88	307	22500 $\pm$ 2250	160	154	12100 $\pm$ 1210
Mrk 463	63	...	2060 $\pm$ 210	88	...	1870 $\pm$ 200	157	...	1300 $\pm$ 140
NGC 2623	64	71	28600 $\pm$ 2860	90	61	26300 $\pm$ 2630	160	49	17200 $\pm$ 1720
NGC 6090	65	209	7000 $\pm$ 700	81	234	8370 $\pm$ 840	163	57	6740 $\pm$ 670
UGC 4881	66	154	5150 $\pm$ 520	92	153	9640 $\pm$ 960	164	69	6680 $\pm$ 670
UGC 5101	63	284	12400 $\pm$ 1240	92	48	17800 $\pm$ 1780	151	122	16400 $\pm$ 1640
VV 705	65	206	9580 $\pm$ 960	82	128	10500 $\pm$ 1050	164	65	5910 $\pm$ 590

## APPENDIX C: AGN OBSERVABLES LINEAR FIT ANALYSIS

**Table C1.** The linear regression test results of  $f_{\text{AGN}}$  versus colours showing the strongest correlations (Pearson  $r > 0.8$ ) and their significance across both subsamples (requiring at least  $3\sigma$  significance).

Flux Ratio	Number of systems	Slope	$r$	$p$
GALEX FUV–MIPS 24 $\mu\text{m}$	20	$6.61 \pm 1.08$	0.82	$8.47 \times 10^{-6}$
GALEX FUV–IRAS 25 $\mu\text{m}$	18	$6.25 \pm 1.03$	0.83	$1.7 \times 10^{-5}$
GALEX FUV–IRAS 60 $\mu\text{m}$	19	$6.13 \pm 0.99$	0.83	$1.03 \times 10^{-5}$
GALEX FUV–PACS 70 $\mu\text{m}$	20	$6.11 \pm 1.05$	0.81	$1.55 \times 10^{-5}$
GALEX FUV–MIPS 70 $\mu\text{m}$	13	$5.36 \pm 1.14$	0.82	$0.06 \times 10^{-2}$
GALEX NUV–IRAS 12 $\mu\text{m}$	20	$3.99 \pm 0.65$	0.82	$7.8 \times 10^{-6}$
GALEX NUV–MIPS 24 $\mu\text{m}$	21	$6.37 \pm 0.88$	0.86	$7.06 \times 10^{-7}$
GALEX NUV–IRAS 25 $\mu\text{m}$	19	$6.07 \pm 0.84$	0.87	$1.41 \times 10^{-6}$
GALEX NUV–IRAS 60 $\mu\text{m}$	20	$5.92 \pm 0.85$	0.85	$1.6 \times 10^{-6}$
GALEX NUV–PACS 70 $\mu\text{m}$	21	$5.79 \pm 0.88$	0.83	$2.67 \times 10^{-6}$
GALEX NUV–MIPS 70 $\mu\text{m}$	14	$5.23 \pm 0.95$	0.85	$1.31 \times 10^{-4}$
Sloan $u$ –MIPS 24 $\mu\text{m}$	23	$6.34 \pm 0.96$	0.82	$1.55 \times 10^{-6}$
Sloan $u$ –MIPS 70 $\mu\text{m}$	15	$4.98 \pm 0.94$	0.83	$1.48 \times 10^{-4}$
Sloan $g$ –MIPS 24 $\mu\text{m}$	23	$6.63 \pm 1.01$	0.82	$1.77 \times 10^{-6}$
Sloan $g$ –PACS 70 $\mu\text{m}$	23	$6.04 \pm 0.97$	0.81	$3.45 \times 10^{-6}$
Sloan $g$ –MIPS 70 $\mu\text{m}$	15	$5.4 \pm 1.02$	0.83	$1 \times 10^{-4}$
Sloan $r$ –MIPS 24 $\mu\text{m}$	23	$6.69 \pm 1.03$	0.82	$2.02 \times 10^{-6}$
Sloan $r$ –MIPS 70 $\mu\text{m}$	15	$5.56 \pm 1.05$	0.83	$1.45 \times 10^{-4}$
Sloan $i$ –MIPS 24 $\mu\text{m}$	23	$6.7 \pm 1.05$	0.81	$2.6 \times 10^{-6}$
Sloan $i$ –MIPS 70 $\mu\text{m}$	15	$5.64 \pm 1.07$	0.83	$0.15 \times 10^{-3}$
Sloan $z$ –MIPS 24 $\mu\text{m}$	23	$7.01 \pm 1.04$	0.83	$1.16 \times 10^{-6}$
Sloan $z$ –PACS 70 $\mu\text{m}$	23	$6.35 \pm 0.99$	0.81	$2.45 \times 10^{-6}$
Sloan $z$ –MIPS 70 $\mu\text{m}$	15	$6.01 \pm 1.04$	0.85	$6.46 \times 10^{-5}$
Sloan $z$ –PACS 100 $\mu\text{m}$	18	$5.16 \pm 0.93$	0.81	$4.56 \times 10^{-5}$
2MASS J–MIPS 24 $\mu\text{m}$	23	$6.98 \pm 1.01$	0.83	$8.44 \times 10^{-7}$
2MASS J–PACS 70 $\mu\text{m}$	23	$6.36 \pm 0.97$	0.82	$1.68 \times 10^{-6}$
2MASS J–MIPS 70 $\mu\text{m}$	15	$6.09 \pm 1.07$	0.85	$7.12 \times 10^{-5}$
2MASS J–IRAS 100 $\mu\text{m}$	12	$8.6 \pm 1.81$	0.83	$7.71 \times 10^{-4}$
2MASS J–PACS 100 $\mu\text{m}$	18	$4.97 \pm 0.93$	0.8	$6.61 \times 10^{-5}$
2MASS H–IRAC 4.5 $\mu\text{m}$	24	$3.85 \pm 0.61$	0.8	$2.52 \times 10^{-6}$
2MASS H–IRAC 5.8 $\mu\text{m}$	24	$4.44 \pm 0.7$	0.8	$2.2 \times 10^{-6}$
2MASS H–MIPS 24 $\mu\text{m}$	23	$6.86 \pm 1.04$	0.82	$1.49 \times 10^{-6}$
2MASS H–PACS 70 $\mu\text{m}$	23	$6.33 \pm 0.94$	0.83	$1.21 \times 10^{-6}$
2MASS H–MIPS 70 $\mu\text{m}$	15	$5.95 \pm 1.14$	0.82	$1.62 \times 10^{-4}$
2MASS H–IRAS 100 $\mu\text{m}$	12	$8.78 \pm 1.7$	0.85	$4.31 \times 10^{-4}$
2MASS H–PACS 100 $\mu\text{m}$	18	$5.12 \pm 0.87$	0.83	$2.26 \times 10^{-5}$
2MASS Ks–PACS 70 $\mu\text{m}$	23	$5.81 \pm 0.87$	0.82	$1.39 \times 10^{-6}$
2MASS Ks–IRAS 100 $\mu\text{m}$	12	$7.96 \pm 1.67$	0.83	$7.79 \times 10^{-4}$
2MASS Ks–PACS 100 $\mu\text{m}$	18	$4.61 \pm 0.76$	0.83	$1.65 \times 10^{-5}$
IRAC 3.6 $\mu\text{m}$ –IRAC 4.5 $\mu\text{m}$	24	$1.38 \pm 0.19$	0.84	$2.89 \times 10^{-7}$
IRAC 3.6 $\mu\text{m}$ –IRAS 100 $\mu\text{m}$	12	$7.07 \pm 1.36$	0.85	$1 \times 10^{-4}$
IRAC 4.5 $\mu\text{m}$ –IRAS 100 $\mu\text{m}$	12	$5.64 \pm 1.21$	0.83	$9 \times 10^{-4}$
IRAC 5.8 $\mu\text{m}$ –IRAS 100 $\mu\text{m}$	12	$3.68 \pm 0.79$	0.83	$8.78 \times 10^{-4}$
IRAS 60 $\mu\text{m}$ –IRAS 100 $\mu\text{m}$	12	$-2.81 \pm 0.46$	$-0.89$	$1.07 \times 10^{-4}$
IRAS 60 $\mu\text{m}$ –MIPS 160 $\mu\text{m}$	14	$-3.25 \pm 0.59$	$-0.85$	$1.4 \times 10^{-4}$
PACS 70 $\mu\text{m}$ –MIPS 160 $\mu\text{m}$	14	$-2.46 \pm 0.44$	$-0.85$	$1.19 \times 10^{-4}$
MIPS 70 $\mu\text{m}$ –IRAS 100 $\mu\text{m}$	7	$-2.13 \pm 0.37$	$-0.93$	$2.23 \times 10^{-3}$
MIPS 70 $\mu\text{m}$ –MIPS 160 $\mu\text{m}$	13	$-2.72 \pm 0.47$	$-0.87$	$1.14 \times 10^{-4}$
MIPS 70 $\mu\text{m}$ –PACS 160 $\mu\text{m}$	14	$-2.74 \pm 0.49$	$-0.85$	$1.3 \times 10^{-4}$
MIPS 70 $\mu\text{m}$ –SPIRE 250 $\mu\text{m}$	15	$-3.4 \pm 0.54$	$-0.87$	$2.56 \times 10^{-5}$
MIPS 70 $\mu\text{m}$ –SPIRE 500 $\mu\text{m}$	15	$-4.23 \pm 0.68$	$-0.87$	$3.08 \times 10^{-5}$
MIPS 70 $\mu\text{m}$ –SPIRE 350 $\mu\text{m}$	15	$-3.75 \pm 0.61$	$-0.86$	$3.46 \times 10^{-5}$
IRAS 100 $\mu\text{m}$ –PACS 160 $\mu\text{m}$	12	$-3.1 \pm 0.35$	$-0.94$	$5.44 \times 10^{-6}$
IRAS 100 $\mu\text{m}$ –SPIRE 500 $\mu\text{m}$	12	$-5.69 \pm 1.03$	$-0.87$	$2.49 \times 10^{-4}$
IRAS 100 $\mu\text{m}$ –SPIRE 350 $\mu\text{m}$	12	$-5.09 \pm 0.81$	$-0.89$	$9.02 \times 10^{-5}$
IRAS 100 $\mu\text{m}$ –SPIRE 250 $\mu\text{m}$	12	$-4.5 \pm 0.65$	$-0.91$	$3.99 \times 10^{-5}$
PACS 100 $\mu\text{m}$ –PACS 160 $\mu\text{m}$	18	$-1.35 \pm 0.25$	$-0.8$	$5.9 \times 10^{-5}$

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